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News & Views

Findings of the Mars Special Regions Science Analysis Group

THE MEPAG SPECIAL REGIONS–SCIENCE ANALYSIS GROUP

EXECUTIVE SUMMARY

Introduction and approach

Current planetary protection (PP) protection policy designates a categorization IVc for spacecraft potentially entering into a "special region" of Mars that requires specific constraints on spacecraft development and operations.

National Aeronautics and Space Administration (NASA) requested that Mars Exploration Program Analysis Group (MEPAG) charter a Special Regions–Science Analysis Group (SR-SAG) to develop a quantitative clarification of the definition of "special region" that can be used to distinguish between regions that are "special" and "non-special" and a preliminary analysis of specific environments that should be considered "special" and "non-special."

The SR-SAG used the following general approach: Clarify the terms in the existing Committee on Space Research (COSPAR) definition; establish temporal and spatial boundary conditions for the analysis; identify applicable threshold conditions for propagation; evaluate the distribution of the identified threshold conditions on Mars; analyze on a case-by-case basis those purported geological environments on Mars that could potentially exceed the biological threshold conditions; and, furthermore, describe conceptually the possibility for spacecraft-induced conditions that could exceed the threshold levels for propagation.

The following represent the results of the SR-SAG study in which "special regions" are more practically defined, including a comprehensive distillation of our current understanding of the limits of terrestrial life and their relationship to relevant martian conditions. An analytical ap-

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proach is presented to consider special regions with current and future improvements in our understanding. The specific findings of the SR-SAG reported in the executive summary are in bold.

Definition

The existing definition of "special region" (from COSPAR, 2005; NASA, 2005) is " . . . a region within which terrestrial organisms are likely to propagate, or a region which is interpreted to have a high potential for the existence of extant martian life forms. Given current understanding, this applies to regions where liquid water is present or may occur." The SR-SAG determined that, to proceed with identifying special regions, some words needed clarification. The word propagate is taken to mean reproduction (not just growth or dispersal). Also, the focus on the word "likely" is taken to apply to the probability of specific geological conditions during a certain time period and not to probability of growth of terrestrial organisms. While the report does concentrate on the salient parameters of forward contamination and martian environmental conditions, it does not address the second clause of the definition concerning probability of martian life, as there is no information.

The study limited itself to special regions that may exist on Mars and to environmental conditions that may exist within the next 100 years, a period reasonably within our predictive capabilities and within which astronauts are expected to be on the surface of Mars. The SR-SAG also considered only the upper 5 m of the Red Planet as the maximum depth that current spacecraft could access as a consequence of failure during entry, descent, and landing. Environments deeper than 5 m were also considered important as possible habitats for life and targets for future exploration. However, in the absence of specific information about the subsurface environment and the operational approach of any future robotic platform to access the deep subsurface, the SR-SAG recommended that such cases should be analyzed on a case-by-case basis.

Limits to microbial life

The approach of the study group was to find any terrestrial representative that demonstrated the ability to reproduce under the worst environmental conditions. Although many factors may limit microbial growth and reproduction, the known overriding environmental constraints on Mars are low temperature and aridity, and a surface that is bathed in ultraviolet (UV) and galactic cosmic radiation.

Life on Earth has been able to survive extremely low temperatures, but for this study, the figure of merit is the ability to reproduce. An extensive review of the literature on low temperature metabolic/reproductive studies reveals that an exponential decrease in microbial metabolism enables long-term survival maintenance or perhaps growth. However, experiments and polar environments themselves have failed to show microbial reproduction at temperatures below −15°C. **For this reason, with margin added, a temperature threshold of −20°C is proposed for use when considering special regions.**

Although many terrestrial microorganisms can survive extreme desiccation, they all share the absolute requirement for liquid water to grow and reproduce. Various measures are used to quantify the availability of liquid water to biological systems, but the one that was used to integrate biology and geology for this analysis was water activity (a_w). Pure water has an a_w of 1.0, and the value decreases with increasing solute concentration and with decreasing relative humidity. Some example a_w values are: seawater = 0.98, saturated NaCl = 0.75, ice at −40°C = 0.67. For this application, a_w has the advantage in that it is a quantity that can be derived and measured, and applied across multiple length scales in equilibrium. The lowest known a_w that allows microbial growth is for a yeast in an 83% (wt/vol) sucrose solution where $a_w = 0.62$. **Based on current knowledge, terrestrial organisms are not known to be able to reproduce at an a_w below 0.62; with margin, an activity threshold of 0.5 is proposed for use when considering special regions.**

Water on Mars

Water on Mars in best analyzed in two broad classifications: the portions of Mars that are at or close to thermodynamic equilibrium and those that are in long-term disequilibrium.

In considering martian equilibrium conditions, the repeatability of thermal inertia results from data set to data set suggests that numerical thermodynamic models are generally accurate to better than a few degrees during most seasons and are even more accurate on an annual average. Comparison between Mars Odyssey Gamma Ray Spectrometer (GRS) measurements and theoretical models of ice stability based on these same thermodynamic numerical models demonstrates excellent agreement between theory and observation. A critically important value of models is that they have predictive value down to spatial scales much finer than that achievable by observational data, and so, though there are macroscopic processes that can produce distinct departures from equilibrium, the scale tends to be local to regional, not microscopic.

Where ice is in vapor-diffusive exchange with the atmosphere, the equilibrium temperature (the frost point) is at about −75°C on contemporary Mars. Ice is not stable with respect to sublimation in places where diurnal or seasonal temperature fluctuations significantly exceed −75°C. Thus, Mars' ample supply of near-surface water is stubbornly sequestered in solid form at temperatures below the frost point, either on the polar caps or in vast high-latitude, subsurface deposits. While the surface of Mars at many low-latitude locations may exceed 0°C in the peak of the day, the temperature 10–20 cm below those surfaces remains perpetually below −40°C. Were liquid to form at a higher surface temperature, it would be transported in a matter of *minutes* or *hours* to the relatively cold region just below the surface, and eventually to a permanent polar or subpolar reservoir by evaporation and condensation. Thus, persistent liquid water at or near the martian surface requires a significant departure from the general planetary setting in the form of either long-term disequilibria (such as geothermal sources) or short-term disequilibria (an impactor).

The equilibrium a_w of martian regolith can be calculated as a function of temperature, using a mean absolute humidity of 0.8 μbar and assuming equilibrium with the atmosphere. In warm regolith, a_w is literally orders of magnitude too small to support life. The a_w approaches unity at the frost point, but at extremely low temperatures. If, however, there is a significant barrier to equilibration with the atmosphere, there is a possibility of much higher absolute humidity and, therefore, significantly higher a_w at warmer temperatures. Desert crusts have been proposed as a potential mechanism to provide a diffusion barrier, and were considered in this study. **Although crusts on Mars have been observed at the past landing sites, and other crust types are hypothetically possible elsewhere, experience with desert crusts on Earth shows that the effect of a semipermeable crust is to retard, not prevent, the achievement of equilibrium.**

Where the surface and shallow subsurface of Mars are at or close to thermodynamic equilibrium with the atmosphere (using time-averaged, rather than instantaneous, equilibrium), temperature and a_w in the martian shallow subsurface are considerably below the threshold conditions for propagation of terrestrial life. The effects of thin films and solute freezing point depression are included within the a_w.

An extensive literature speculates on mechanisms to form liquid water on Mars at different times in the past and under different climate conditions (e.g., Farmer, 1976; Clow, 1987; Carr, 1996), and common to all of them is the explicit understanding that present-day equilibrium conditions do not support the persistence of liquid water at the surface. Uncertainty exists as to whether previous conditions were persistent or episodic, with some attributing conditions to be punctuated, due to impact effects, and others envisioning longer-term stable early climates. More recently, orbital forcing has been recognized as a factor that drives climate change, with 50,000 years being the shortest climate cycle affecting latitudinal precipitation.

The SR-SAG considered possible environments in long-term disequilibrium, where water and temperature were in equilibrium under conditions at an earlier time, but for which conditions have changed, and do not hold for the present. Geological deposits might survive for 10^4–10^7 years by virtue of giving up their water very slowly. The SR-SAG examined several potential sites for long-term disequilibrium, either theoretical or actually observed, such as gullies, mid-latitude features of purported snow/ice deposits, remnant glacial deposits, craters, volcanoes, slope streaks, recent outflow channels, possible hydrothermal vents, low-latitude ground ice, and polar caps.

- **Some—though certainly not all—gullies and gully-forming regions might be sites at which liquid water comes to the surface within the next 100 years. At present, there are no known criteria by which a prediction can be made as to which—if any—of the tens of thousands of gullies on Mars could become active—and whether the fluid involved is indeed water—during this century.**

- **Because some of the "pasted-on"–type mantle has a spatial, and possibly a genetic, relationship to gullies (which in turn are erosional features possibly related to water), the "pasted-on" mantle may be a special region. The mid-latitude mantle, however, is thought to be desiccated, with low potential for the possibility of transient liquid water in modern times. Because the "pasted-on" mantle and some kinds of gullies may have a genetic relationship, the mantle is interpreted to have a significant potential for modern liquid water.**
- **No craters with the combination of size and youthfulness to retain enough heat to exceed the temperature threshold for propagation have been identified on Mars to date.**
- **We do not have evidence for volcanic rocks on Mars of an age young enough to retain enough heat to qualify as a modern special region or suggest a place of modern volcanic or hydrothermal activity.**
- **Despite a deliberate and systematic search spanning several years, no evidence has been found for the existence of thermal anomalies capable of producing near-surface liquid water.**
- **The martian polar caps are too cold to be naturally occurring special regions in the present orientation of the planet.**

The SR-SAG proposes that martian regions may be categorized as non-special if the temperature will remain below -20°C *or* the a_w will remain below 0.5 for a period of 100 years after spacecraft arrival. All other regions on Mars are designated as either special or uncertain. An uncertain region is treated as special until it is shown to be otherwise. The SR-SAG found no regions to be special, but found uncertainty with the gully and possibly related "pasted-on" mantle regions. In this context, the SR-SAG has listed Mars environments that may be "special" and classified those that have observed features probably associated with water, those that have a nonzero probability of being associated with water, and those areas that, if found, would have a high probability of being associated with water.

A map has been developed that provides generalized guidelines for the distribution of areas of concern that may be treated as special regions.

It should be noted that, even in a region determined to be "non-special," it is possible that a spacecraft could create an environment that meets the definition of "special" or "uncertain." It is possible for spacecraft to induce conditions that could exceed for some time the threshold conditions for biological propagation, even when the ambient conditions were "not special" before the spacecraft arrived. Whether a special region is induced or not depends on the configuration of the spacecraft, where it is sent, and what it does. This possibility is best evaluated on a case-by-case basis.

In summary, within the upper 5 m most of Mars is either too cold or too dry to support the propagation of terrestrial life. However, there are regions that are in disequilibrium, naturally or induced, and could be classified as "special" or, if enough uncertainty exists, could not be declared as "non-special."

Key Words: • Mars • Extremophile microorganism • Habitability • Planetary protection • Water activity • Special region

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1. INTRODUCTION

1A. "Special Regions"—History and the Current Problem

In 2002, COSPAR introduced the term "special region" as a part of Mars PP policy. Prior to 2002, PP-related requirements for spacecraft going to the martian surface consisted of two categories that were distinguished by the purpose of the mission:

- IVa. Landers without extant life detection investigations
- IVb. Landers with extant life detection investigations

By 2002, however, exploration results [primarily from the Mars Global Surveyor (MGS) orbiter, and soon after confirmed by Mars Odyssey] strongly suggested that some parts of Mars might be more likely than others to attract interest for extant life investigations and, potentially, more vulnerable to the effects of Earth-sourced biological contamination. This led to the introduction of the concept of "special regions," which are environments on Mars that need a high degree of protection independent of the mission purpose.

In April 2002, a COSPAR planetary protection workshop formulated a draft definition of "special region" and proposed that a new mission categorization, Category IVc, be established for missions that come (or might come) into contact with them. This proposal was presented to COSPAR at its 2002 meeting, and was formally adopted shortly afterwards (http://www.cosparhq.org/scistr/PPPolicy.htm). NASA followed up by incorporating the special regions concept into its policy by means of modification of NASA Procedural Requirements 8020.12C *Planetary Protection Provisions for Robotic Extraterrestrial Missions*, which was issued in 2005.

In 2005, a National Research Council (NRC) committee (referred to as NRC PREVCOM) completed a NASA-requested detailed 2-year study entitled *Preventing the Forward Contamination of Mars* (NRC, 2006). [NRC PREVCOM was a committee of the National Research Council (of the National Academies of Science) that, at NASA's request, examined PP measures for Mars. Subsequent to accepting its statement of task, an NRC committee operates independently of its sponsoring agency.] In their analysis of "special regions," NRC PREVCOM found that, in using the current special region definition, "there is at this time insufficient data to distinguish with confidence "special regions" from regions that are not special." They also raised an important issue of scale—"Mars exhibits significant horizontal and spatial diversity on km to cm spatial scales," but some of the relevant observational data have a spatial resolution no better than $\sim 3 \times 10^5$ km^2. NRC PREVCOM recommended an interim policy in which all of Mars is considered a "special region."

For further information on PP policy and history related to Mars, the interested reader is referred to excellent recent reviews by DeVincenzi *et al*. (1998) and NRC (2006).

> ***DEFINITION #1.***
>
> Existing definition of "special region" (from COSPAR, 2005; NASA, 2005):
>
> " . . . a region within which terrestrial organisms are likely to propate, or a region which is interpreted to have a high potential for the existence of extant martian life forms. Given current understanding, this applies to regions where liquid water is present or may occur. Specific examples include but are not limited to:
> a) Subsurface access in an area and to a depth where the presence of liquid water is probable
> b) Penetration into the polar caps
> c) Areas of hydrothermal activity"

1B. This Study

Purpose. At the November 2005 MEPAG meeting, NASA requested that MEPAG prepare a community-based analysis of the definition of "special region" and, if possible, propose clarifications that make the definition more useful for mission planning and PP implementation. MEPAG in turn chartered the SR-SAG and gave it the following assignment:

- Propose, if it is possible to reach consensus, a quantitative clarification of the definition of "special region" that can be used in a practical way to distinguish between regions on Mars that are "special," "non-special," and "uncertain."
- Prepare a preliminary analysis, in text form, of the kinds of martian environments that should

be considered "special" and "non-special." If possible, also represent this in map form.

Methodology. The SR-SAG consisted of 27 members with scientific backgrounds in various aspects of microbial survival, physics, geology, and PP. The group included three members who also served as part of NRC PREVCOM. The SAG met by means of weekly teleconferences (with several subgroups working in parallel) in December 2005 and January 2006, along with extensive e-mail exchange. From February 6 to 8, a 3-day Special Regions Workshop was held in Long Beach, CA, to integrate results.

1C. How Does This Study Extend the Results of PREVCOM?

We consider the present study to be an extension of the work of the NRC's PREVCOM Committee (NRC, 2006). Given the phrasing of COSPAR's definition of special regions and, more importantly, the "specific examples" listed, NRC PREVCOM brought forward their recommendation that "until measurements are made that permit confident distinctions to be drawn between regions that are special on Mars and those that are not, NASA should treat all direct contact missions as category IVs" [missions to special regions, for which they recommended specific biological cleanliness requirements]. NRC PREVCOM worked with the existing definition and elected not to recommend modifications or qualifications to COSPAR's language. [The NRC PREVCOM's Statement of Task included the language that "to the maximum possible extent, the recommendations should be developed to be compatible with an implementation that would use the regulatory framework for planetary protection currently in use by NASA and the Committee on Space Research (COSPAR)." The full NRC PREVCOM statement of task is given on pp. vii–viii, NRC PREVCOM (2006).] They advised that the community should endeavor to expand current understanding through measurement and analysis in order "to permit confident distinctions to be drawn." This led to the purpose of the SR-SAG, which was to consider the COSPAR definition and propose necessary and appropriate clarifications, qualifications, and extensions that would allow an improved ability to recognize special regions (and to allow different people to reach the same interpretation of the definition).

NRC PREVCOM was explicit in its advice that the Mars Program should pursue measurements to define special regions. While this study recognizes that models carry uncertainty and measurements will be forthcoming in the course of exploration, we have extended currently available information through the use of very conservative models and analysis.

1D. Future Steps

Our knowledge about Mars and the limits of life on Earth will continue to evolve in the coming years. While the analysis reported here has attempted to make conservative assumptions and add additional margins to proposed thresholds, the SR-SAG anticipates that findings reported here may be reviewed and, if necessary, updated several years from now unless sudden discoveries require an earlier revision.

2. APPROACH

The charge to the SR-SAG was to prepare a community-based analysis of the definition of "special region" and propose clarifications and/or guidelines that make the definition less ambiguous and more practical. The SR-SAG used the following general approach:

1. Consider the terms in the existing COSPAR definition and clarify as needed.
2. Establish temporal and spatial boundary conditions for analysis.
3. Identify applicable threshold conditions for propagation of terrestrial organisms.
4. Evaluate the distribution of the identified threshold conditions on Mars, using both data and models, as appropriate.
5. Analyze on a case-by-case basis those geological environments (including those that are hypothetical) on Mars that could (or would if they existed) potentially exceed the biological threshold conditions.
6. Describe conceptually the possibility for spacecraft-induced conditions that could exceed the threshold levels for propagation; and propose an approach to respond to this possibility.

A comment about the scientific literature pertaining to water on Mars. There is a very large, and what appears at first glance to be conflicting, literature relating to water on Mars. This has created a certain confusion in the community. However, the conclusions of many of the papers in the literature have qualifications that involve time or circumstances. To facilitate interpretation of the literature and the application of it to specifics of the special region question, the SR-SAG found it necessary to start from first principles to derive its own understanding of the potential for water on Mars during the time period of interest. This has given SR-SAG a context for assimilating and integrating the many relevant details in the literature.

3. CLARIFICATION OF THE EXISTING SPECIAL REGION DEFINITION

The special region definition (above, DEFINITION #1) consists of two parts: (1) a defining statement that consists of two clauses and (2) a description of where, under the current interpretation, special regions may occur. The SR-SAG concludes that the first part is still useful, as long as some of the terms are clarified. The second half needs to be revised and extended with an updated statement of "current understanding."

The first clause of the defining statement includes the following words, which need clarification:

- *Propagate*. The verb "propagate" has two meanings, for which the respective synonyms are "reproduce" and "spread." For the purpose of this analysis, we have assumed the former meaning only. Although there has been extensive discussion that a biological contamination event requires *both* reproduction and dispersion to create a problem for future explorers, a more conservative position is that reproduction alone is sufficient to create questions, and this was taken as the point of departure for this study.
- *Likely*. It is assumed for the purpose of this analysis that the probability of growth of terrestrial organisms under all martian environmental conditions cannot be accurately determined. However, the probability that specified geological conditions exist within a certain time period can be estimated, in some cases quantitatively.

The second clause in the defining statement pertains to possible martian life forms and their likely locations. Because there is no information on martian life forms, the hardiest Earth organisms are used as a proxy. However, the clause remains as part of the definition since, in the future, our understanding of potential martian life may change and affect the parameters that define special regions. As a consequence, the SR-SAG analysis and this report concentrate on the forward contamination of Mars with live organisms from Earth. The focus here is on identification of parts of the martian environment in which viable terrestrial organisms would be unable to propagate, and establishment of an objective description of such areas so that appropriate planning and implementation for PP can occur.

4. BOUNDARIES FOR THE PRESENT ANALYSIS

The analysis of martian special regions required certain boundary conditions to be established as a basis for study. One significant boundary condition was the time frame to be used in the identification of special regions. Another was a spatial boundary (depth) to be applied to this analysis. Discussion of these two key boundaries—time and depth—is presented below.

4A. Time frame

With respect to special regions, timeframe issues can be viewed in three ways—how long to avoid special regions, how long do special regions exist, or how long until they may exist. Current PP standards proscribe atmospheric entry by any Mars orbiter for a 50-year period if spacecraft assembly has not incorporated explicit protocols for bioburden reduction beyond assembly in a class 100,000 cleanroom. This time span was selected toward the beginning of Mars exploration, when it was envisioned that the pace of Mars exploration would be quicker than it has been. Because of the technical challenges of accomplishing successful Mars missions, their high cost, and the transition from a "space race" to the more measured pace of international space cooperation, fewer than 20 missions have been launched, and only about a third of those were successfully implemented in the 3 decades since the early Viking missions. Furthermore, from recent orbiter and rover missions it has become recognized that Mars is far more diverse than earlier explorations had indicated, with a very large number of scientific sites now identified for future exploration. Many cognizant researchers now anticipate that the period of biological exploration will span the current century, and this study makes no explicit assumptions about the length of the exploration period.

Based on input from the NASA Planetary Protection Officer, this study used a 100-year time frame over which the existence of martian special regions would be considered and could be encountered by any given mission. This figure was accepted as a premise for the SR-SAG analysis. It allowed for the analysis of martian environments to take into account past and present climate, but not to extend to the distant future of climate change driven by obliquity cycles on Mars. It included consideration of current naturally occurring special regions, the possibility that a region could become a special region within the next 100 years (from the date of a mission's arrival) due to a natural event (*e.g.*, eruption of a volcano), and the time scale for spacecraft-induced special regions.

How might the environmental conditions on Mars over approximately the next 100–200 years differ from those of today? The primary factor that controls long-term climate change on Mars is the variation in the planetary obliquity (the tilt of its spin axis with respect to its orbital plane) with time. The martian obliquity has varied between 15° and 35° during the last 5 million years, with a periodicity of about 120,000 years (Laskar *et al*., 2004). This variation is widely regarded to have been responsible for major climate variations in the past (*e.g.*, Jakosky and Carr, 1985; Haberle *et al*., 2003; Head *et al*., 2003a, 2005, 2006a,b; Mischna *et al*., 2003; Mischna and Richardson, 2005; Forget *et al*., 2006). For example, when the obliquity is greater than about 30°, the annually averaged saturation vapor pressure at the martian poles is greater than at the equator, a condition that drives a major redistribution of both water and CO_2 on a planetary scale. At present, Mars has a tilt of 25.2° and is about halfway through one of these obliquity cycles, though it is presently in a quiescent period of very little obliquity change. This means that 100 years from now the martian obliquity will be only marginally higher than at present, which is not of significance for long-term climate change (Nakamura and Tajika, 2003).

The south polar cap does appear to be able to change within a 100-year time scale. There are observations that show changes in the CO_2 ice cover from one year to another (Malin *et al*., 2001; Thomas *et al*., 2005) and changes on the decade time scale in the outline of the cap (or equivalently the degree of CO_2 ice cover). Observations have also shown that water ice is exposed where the CO_2 ice is disappearing (Titus *et al*., 2003; Bibring *et al*., 2004). In addition, there are less direct inferences from the water vapor seasonal behavior over many decades that have suggested the same type of behavior but possibly with more extreme results (*e.g*., the entire CO_2 cap potentially disappearing in some years). From a stability standpoint, there is no reason why the CO_2 ice cannot come and go, possibly on the decade to century time scale (see Jakosky *et al*., 2005a,b). Whether and how this might affect climatic conditions elsewhere on Mars is not known. However, we do not have evidence that these south polar CO_2 effects are causing significant changes in the planetary distribution of water.

The SR-SAG consensus is that the martian climate 100 years (and 1,000 years) from now will likely be essentially the same as it is today.

PREMISE. A 100-year time span may be used to assess the potential for special regions that may be encountered by any given mission.

4B. Maximum Depth of Penetration by an Impacting Spacecraft

While PP concerns itself with all of Mars (surface and subsurface), not all of Mars is accessible to contamination by robotic spacecraft. Thus, a practical analysis of special regions must take into consideration the part of the surface and shallow subsurface that is vulnerable to contamination. For all missions, aside from planned operations, there is the possibility of accidental subsurface access as a result of hard impact (*i.e*., a crash). There can also be access to the subsurface as a result of intentional hard impact (*e.g*., end of mission disposal of hardware or hard landing of entry, descent, and landing hardware). To address these issues, it is possible to analyze impact scenarios and physical conditions at Mars to put bounds on the possible contamination depth.

The depth of penetration of a crashing spacecraft is a function of the following parameters: the angle of impact, the impact velocity, the mass of the impacting object, and the strength and density of the geological material being impacted. All of these parameters will vary from mission to mission. The impact velocity is dependent on the entry velocity (at the top of the atmosphere) and the ballistic coefficient, which determines how much the spacecraft will be slowed by the martian atmosphere. Spacecraft sent to Mars in the future will have a range of ballistic coefficients, and entry velocity will be different for each launch opportunity and will also depend on the choice of trajectory. The penetration depth depends on whether the mass of the spacecraft stays together or breaks up as it passes through the atmosphere. The impact angle in a failure scenario would depend on when control of the trajectory were lost. Finally, the martian surface consists of a mixture of outcrop (of both igneous and sedimentary rocks), regolith, accumulated wind-blown dust, and polar cap material, all of which could have been cemented by ice and/or minerals and would influence the penetration depth.

The depth of impact can be estimated with crater scaling laws. For impact into dry granular regolith, the following fit to dry sand impact data is suitable (Holsapple, 1993):

$$V = 0.14\ (1700/\rho)\ M^{0.83}\ U^{1.02}/G^{0.51}$$

where V is crater volume in m^3, ρ is the regolith density in kg/m^3, M is the impacting mass in kg, U is the velocity in km/s, and G is the strength of gravity relative to Earth (about 0.38). The term $(1700/\rho)$ has been included to extend the original model to densities other than the nominal sand density of 1,700 kg/m^3.

For impacts into icy material the following weak rock fit is used (Holsapple, 1993):

$$V = 0.009\ (2100/\rho)\ M\ U^{1.65}$$

Again, the model has been extended with a density dependence. This model is intended for impacts into targets with strengths averaging about 7.6 MPa over large areas. The laboratory strength of frozen soils and ice are on the order of 20 MPa at −25°C (Lee *et al.*, 2002), and higher at lower temperatures, so even allowing for a reduction in strength due to size effects, this model may overestimate crater sizes to some extent. This is appropriate for the purpose of estimating maximum depths.

It remains to specify how the crater depth and diameter are related to the volume. The assumption will be made that the crater is a paraboloid with a depth-to-diameter ratio of 1/4. This is a typical ratio for the maximum transient dimensions of a simple crater. It should approximately agree with the final ratio of an icy crater, but the final crater in dry granular material would be shallower. The volume of a paraboloid with depth H and diameter D is $V = \pi\ H\ D^2/8$, and the assumption $H/D = 1/4$ leads to the depth being $H = 0.54\ V^{1/3}$.

It would be possible to assume a worst-case scenario for each of the above variables for a hypothetical fleet of future spacecraft and, from that, to estimate the maximum theoretical crater depth. However, this would entail a set of stacked probabilities for which the single worst outcome lacks practicality and usefulness. Because of the broad range of possible mission scenarios, rather than attempting to seek out the theoretical maximum, a population of calculated solutions is shown in Figs. 1 and 2. These diagrams assume a perpendicular impact angle (the worst case for that variable) and show some of the relationships that involve impact velocity, mass, and target geology on crater depth. The upper curve in Fig. 1 represents the case for impact into dry regolith material with an extremely low average density of 1,100 kg/m^3.

A relevant scenario is the case of a spacecraft launched on a modern heavy launch vehicle having a mass for the entry system of about 2,400 kg, for which the mass passes intact through the atmosphere and impacts the surface with a velocity of about 4 km/s. Such a system could create a crater with a depth of about 5 m. For other mission scenarios, these diagrams can be used to estimate the possibility of penetrations deeper than

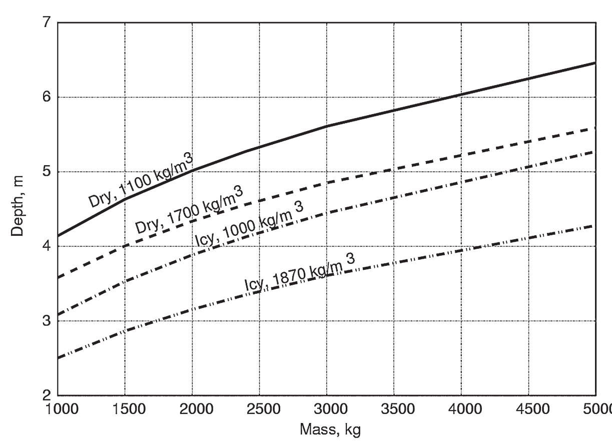


FIG. 1. Crater depth for a spacecraft impacting Mars at 4 km/s (four regolith characteristics shown).

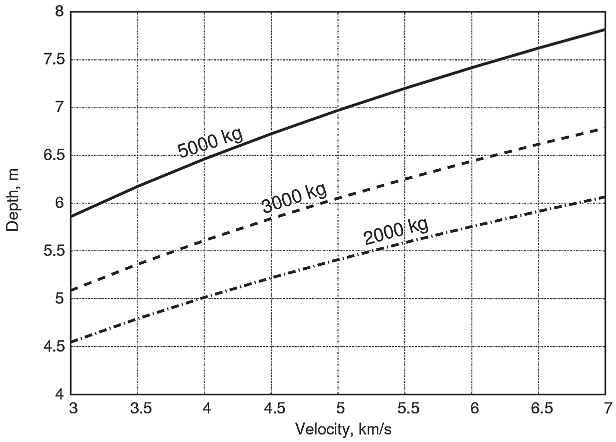


FIG. 2. Crater depth for a spacecraft impact into dry regolith of density 1,100 kg/m^3 over a range of impact velocities (three spacecraft masses shown).

5 m. For example, a hypothetical 5,000 kg spacecraft (larger than we can currently land at Mars) impacting at 4 km/s would have an estimated maximum penetration depth of about 6.5 m.

In the future, we can expect innovative mission concepts to incorporate deliberate access of the deep subsurface through hard impacts, innovative drills, or melt probes. For these, it will be necessary to analyze the possibility of deliberate access into naturally occurring special regions as a result of planned exploration into the deeper martian subsurface. In addition, entry systems at some time in the future will certainly be configured with different masses and ballistic coefficients (*e.g.*, to fit in the launch vehicle fairing), or might arrive at Mars on trajectories with higher atmospheric entry velocities. For these systems, either detailed analysis of atmospheric deceleration can be performed or conservative simplifying assumptions can be used (*e.g.*, no atmosphere) to evaluate impact scenarios and possible consequences.

<u>FINDING</u>. **Although naturally occurring special regions anywhere in the three-dimensional volume of Mars need protection, only those in the outermost ~5 m of the martian crust can be inadvertently contaminated by a spacecraft crash—special regions deeper than that are not of practical relevance for missions with a mass up to about 2,400 kg and possible impact velocities up to ~4 km/s.**

5. IMPLICATIONS FROM MICROBIOLOGY

5A. Introduction

There are many environmental factors to be considered in assessing the ability of microbial life to grow and reproduce (Table 1). As a starting point for our analysis we considered terrestrial life forms that might be capable of growth under extreme conditions of the martian environment, thresholds for environmental factors that would prevent growth and replication, and the physiological and nutritional constraints terrestrial microbes must overcome to pose a threat of widespread forward contamination of Mars over a defined time frame. In general, our strat-

Table 1. Some Factors That May Affect the Survival and Reproduction of Earth Microbes on Mars

Water availability and activity
- Activity of liquid water
- Past/future liquid (ice) inventories
- Salinity, pH, and E_h of available water

Chemical environment
- Nutrients
  - C, H, N, O, P, S, essential metals, essential micronutrients
  - Fixed nitrogen
  - Availability/mineralogy
- Toxin abundances and lethality
  - Heavy metals (*e.g.*, Zn, Ni, Cu, Cr, As, Cd, etc., some essential, but toxic at high levels)
  - Globally distributed oxidizing soils

Energy for metabolism
- Solar (surface and near-surface only)
- Geochemical (subsurface)
  - Oxidants
  - Reductants
  - Redox gradients

Conducive physical conditions
- Temperature
- Extreme diurnal temperature fluctuations
- Low pressure (is there a low-pressure threshold for terrestrial anaerobes?)
- Strong biocidal UVC irradiation
- Galactic cosmic rays and solar particle events (long-term accumulated effects)
- Solar UV-induced volatile oxidants, *e.g.*, O_2^-, O^-, H_2O_2, O_3
- Climate/variability (geography, seasons, diurnal, and, eventually, obliquity variations)
- Substrate (soil processes, rock microenvironments, dust composition, shielding)
- High CO_2 concentrations in the global atmosphere
- Transport (aeolian, ground water flow, surface water, glacial)

Modified after Rummel (2006).

egy has been to find any terrestrial representative (no matter where it is from) that demonstrates the worst-case scenario. We are not assigning any special Mars or spacecraft relevance to any of these organisms or situations, though we are documenting observations that suggest the metabolic or physiological possibility of reproduction.

The Mars environment is extremely cold and dry, and the surface is bathed in UV radiation during the daytime and significantly influenced by galactic cosmic radiation at all times. Because Mars is cold, but not always, and extremely dry, but perhaps not everywhere, the concept of "special region" describes those places where environmental conditions might be compatible with microbial propagation. The special-region concept allows mission planners to address the requirements of PP in regions on Mars where terrestrial Earth organisms might survive and proliferate.

5B. Lower Temperature Threshold

It is well documented that microorganisms on Earth live at temperatures well below the freezing point of pure water, *e.g.*, inside glacial and sea ice and permafrost. This is possible because certain impurities such as mineral acids or salts can reduce the freezing point of water. These impurities can prevent freezing of intergranular veins in ice and thin films in permafrost, and permit transport of nutrients to and waste products from microbes. Furthermore, from viability and survival studies, we know that some cells can resist freezing. Survival strategies include synthesis of stress proteins, reduction in cell size, dormancy, sporulation, adaptive modifications to their cellular components (*e.g.*, changes in their fatty acid and phospholipids composition), or an alteration in the "structured" water in their cytoplasm (Russell, 1992; Thieringer *et al.*, 1998). These and other adaptations allow them to operate more efficiently than mesophilic organisms at low temperatures. Temperature influences growth rates and cell replication by affecting the conformation of cellular macromolecules and other cellular constituents, which in turn control substrate acquisition and determine the rates of enzymes reactions and metabolism (Russell *et al.*, 1990). The relationship between temperature and reaction rate (k) can be described by an Arrhenius equation:

$$k = Ae^{-Ea/RT}$$

where Ea is the activation energy, A is a constant, R is the universal gas constant, and T is absolute temperature. The activation energy for most enzymes is usually on the order of 420 kJ/mol. Therefore, although reactions rates would fall considerably with a drop in temperature, there is no thermodynamic restriction on growth at low temperatures. Although thermodynamics predicts some metabolic activity at low temperatures, the lower temperature limit for cell division is probably set by freezing of the internal solution of the cell rather than reduction in enzymatic activity at low temperature. Therefore, we chose an empirical rather than theoretical approach to setting a lower temperature limit to cell replication.

In developing a rationale for setting a lower temperature threshold, we evaluated published reports of microbial activity that provide direct and/or indirect evidence that microorganisms survive or thrive at temperatures below −5°C. The studies we evaluated fell into three groups: direct measurements of cell replication, measurements of metabolic activity, and indirect measurements of inferred microbial activity (*e.g.*, N_2O production in ice cores). Based on a proposal by Morita (1997), metabolic studies were categorized further into those providing evidence of (1) survival metabolism, *i.e.*, the extremely weak metabolism of immobile, probably dormant communities; (2) maintenance metabolism of communities with access to nutrients, which are free to move but are still below thresholds for growth; or (3) actual growth and cell division that leads to propagation. The metabolic activity measured, the methods used, the temperature limits, and the categories of the responses are listed in Table 2. In addition, several studies have inferred microbial activity below −20°C from anomalous concentrations or stable isotope signatures of products of microbial metabolism. For example, Sowers (2001) proposed nitrification as the likely explanation for peak concentrations of N_2O and high $\delta^{15}N$ and low $\delta^{18}O$ of N_2O in Lake Vostok ice core from the penultimate glacial maximum, about 140,000 years ago. Price and Sowers (2004) estimated that the rates of biomass turnover at −40°C correspond to 10 turnovers of cellular carbon per billion years. Table 2 is not exhaustive, but is representative of a broad and diverse literature on biological activity at low temperatures.

To summarize these data, many groups have demonstrated some metabolic activity (using various measures and by various techniques) at tem-

TABLE 2. OBSERVATIONS OF BIOLOGICAL ACTIVITY AT LOW TEMPERATURES

Reference	*Measurement*	*Temperature minimum*	*Metabolic category*
Bakermans *et al.* (2003)	Cell counts of bacteria isolated from Siberian permafrost	−10°C	Cell replication, DT 39 days
Breezee *et al.* (2004)	*Psychromonas ingrahamii* cell counts from sea ice from off Point Barrow, Alaska	−12°C	Cell replication, DT 10 days
Jakosky *et al.* (2003)	Cell counts of bacteria isolated from Siberian permafrost	−10°C	Cell replication, DT 40 days
Christner (2002)	DNA and protein synthesis by uptake of [^{3}H]thymidine and [^{3}H]leucine, respectively, in psychrotrophs from polar ice cores	−15°C	Maintenance
Gilichinsky *et al.* (2003)	Assimilation of [^{14}C]glucose by bacteria in cryopegs (brine lenses) found in Siberian permafrost	−15°C	Maintenance
Junge *et al*. (2004)	Respiration observed in brine channel prokaryotes in Arctic sea ice communities by CTC	−20°C	Survival
Junge *et al*. (2006)	Protein synthesis, [^{3}H]leucine incorporation	−20°C	Maintenance
Kappen *et al.* (1996)	CO_2 exchange both uptake and loss by polar lichens	−12°C to −18°C	Survival/maintenance?
Rivkina *et al.* (2000)	Incorporation of [^{14}C]acetate into glycolipids by bacterial community from Siberian permafrost	−20°C	Maintenance/replication?, DT 160 days at −10°C?
Rivkina *et al.* (2002)	Measured evolution of methane by a community of permafrost methanogenic archaea	−16.5°C	Survival?
Wells and Deming (2006)	Viral infectivity and production in natural winter sea-ice brines in the Arctic	−12°C	Microbial evolution (lateral gene transfer) and community succession
Carpenter *et al.* (2000)	DNA and protein synthesis by uptake of [^{3}H]thymidine and [^{3}H]leucine, respectively, in psychrotrophs from polar snow	−12°C to −17°C	Maintenance

DT, doubling time; CTC, 5-cyano-2,3-ditoyl tetrazolium chloride.

peratures down to −20°C. At the lowest temperatures, activity was very low (insufficient to support cell replication) and was not sustained beyond a few weeks. Although reported levels of metabolic activity at temperatures down to −15°C might support growth, no one has demonstrated cell replication to occur at or below −15°C. There are no studies that have systematically looked at growth and replication at 1° increments below −15°C. We therefore recommend a lower temperature threshold of −20°C, below which there is no evidence to indicate that replication is possible. (If Earth organisms were to be discovered in the future that were able to replicate at temperatures at or below −20°C, this finding would be reevaluated.)

<u>FINDING</u>. **Based on current knowledge, terrestrial microorganisms are not known to be able to reproduce at a temperature below about −15°C. For this reason, with margin added, a temperature threshold of −20°C is proposed for use when considering special regions.**

5C. Water Activity Threshold

Although many terrestrial microorganisms can survive extreme desiccation in a quiescent state, *e.g.*, as spores, they all share an absolute requirement for liquid water in order to grow, *i.e.*, to multiply and to increase their biomass. Various measures are used to quantify the availability of liquid water to biological systems, depending on the scientific discipline (*e.g.*, soil microbiology, food microbiology, plant physiology, plant pathology). The a_w (that is, the activity of *liquid water*) is related to percent relative humidity (*rh*) as follows:

$$a_w = rh/100$$

when the relative humidity of an atmosphere is in equilibrium with the water in a system (a solution, a porous medium, etc.). For pure water, $a_w = 1.0$. The a_w decreases with increasing concentrations of solutes and as increasing proportions of the water in a system are sorbed to surfaces, *e.g.*, during desiccation in a porous medium such as the martian regolith (Table 3).

Desiccation (matric-induced a_w) and solutes impose related but different stresses on microbial cells. (Matric effects are those induced by the adhesive and cohesive properties of water in contact with a solid surface.) Cytoplasmic a_w must approximate extracellular a_w to avoid excessive turgor (osmotic) pressure, plasmolysis, or plasmoptysis (cell explosion); however, some positive turgor pressure is required for cellular expansion during growth. Microbes respond to decreasing a_w by accumulating intracellular compatible solutes, a response that has been well characterized in many different microorganisms and requires expenditure of energy for transport or synthesis (Brown, 1976, 1990; Csonka, 1989; Welsh, 2000).

Low a_w in a porous medium has the added effect of decreasing nutrient availability. As a soil loses water, the water films on the surfaces of soil particles become thinner and also discontinuous. This limits solute diffusion and also impedes microbial motility. Solute diffusion is reduced by a factor of approximately 2, and microbial mobility is negligible when a soil loses moisture such that a_w drops ~0.99 or less (Wong and Griffin, 1976; Papendick and Campbell, 1981). Thus, low matric-induced a_w in a porous medium imposes starvation conditions due to the diminished solute diffusion and microbial motility. Filamentous organisms (fungi, algae, cyanobacteria, and actinomycetes) may overcome this limitation by extending filaments through air voids in a partially desiccated soil, but this extends their desiccation tolerance only to a_w of approximately 0.9. In the absence of exogenous energy sources, bacteria might be able to undergo two or three rounds of reductive cell division, but this is not an increase in biomass and, thus, is not true growth. Desiccation stress is usually more inhibitory to microbial growth and activity than a solute-induced water stress with an equivalent a_w, primarily because of desiccation-induced nutrient limitation. However,

TABLE 3. CONDITIONS RESULTING IN VARIOUS a_w VALUES AND MICROBIAL RESPONSES TO a_w VALUES

Water Activity (a_w)	*Condition or response*
1.0	Pure water
Solute-induced effects	
0.98	Seawater
0.75	Saturated NaCl solution
0.29	Saturated $CaCl_2$ solution
0.98–0.91	Lower solute-induced a_w limit for growth of various plant pathogenic fungi
0.69	Lower solute-induced a_w limit for growth of *Rhizopus, Chaetomium, Aspergillus, Penicillium* (filamentous fungi)
0.62	Lower solute-induced a_w limit for growth of *Xeromyces* (Ascomycete fungus) and *Saccharomyces* (Ascomycete yeast) (growth in 83% sucrose solution)
Matric-induced effects	
0.999	Average water film thickness = 4 μm
0.9993	Average water film thickness = 1.5 μm
0.996	Average water film thickness = 0.5 μm
0.99	Average water film thickness = 3 nm
0.97	Average water film thickness <3 nm (<10 H_2O molecules thick)
0.93	Average water film thickness <1.5 nm (<5 H_2O molecules thick)
0.75	Average water film thickness <0.9 nm (<3 H_2O molecules thick)
0.999	Matric-induced a_w at which microbial motility ceases in a porous medium
0.97–0.95	Lower matric-induced a_w limit for growth of *Bacillus* spp.
0.88	Lower matric-induced a_w limit for growth of *Arthrobacter* spp.
0.93–0.86	Matric-induced a_w at which microbial respiration becomes negligible in soil

Compiled from Papendick and Campbell (1981), Harris (1981), Griffin (1981), Sommers *et al*. (1981), and Potts (1994).

specific solutes may be toxic to microbes, *e.g.*, sodium ions are inhibitory to some degree to all microbes if they accumulate intracellularly.

There is no doubt that the majority of hypersaline environments on Earth harbor significant populations of microorganisms (for a recent summary, see Grant, 2004). However, values of a_w do not generally fall much below 0.75, the limiting value obtainable at the saturation point of NaCl (5.2 *M*). Halophilic microbes (including members of the Bacteria, Archaea, and Eukarya) can unquestionably propagate in saturated NaCl solutions (a_w = 0.75). Although the presence of organisms in concentrated brines of other salts with a_w lower than 0.75 has been observed, there are questions with regard to the nature of their life cycles and where and how they reproduce and grow.

For example, microbial communities have been reported in Don Juan Pond in Antarctica, a small unfrozen Antarctic lake dominated by very large concentrations of $CaCl_2$ during the winter. Total dissolved salts may exceed 47% (wt/vol), and the a_w value is recorded at 0.45 (Siegel *et al.*, 1979). However, there has been dispute over the evidence for microbial colonization of this site (Horowitz *et al.*, 1972), and the prevailing opinion is that life is unlikely to exist at this a_w value (Grant, 2004). The algal mat communities develop during the summer in melt water at the margins of the pond, which is essentially fresh water, and how this community relates to the low-activity winter brine is uncertain. As summarized by Grant (2004), "this particular site is long overdue for a re-examination using direct molecular technologies." Another example is the $MgCl_2$ and KCl-rich Dead Sea brine (a_w ~0.67). However, the microbes in this brine are likely survivors from brief intervals of growth that follow dilution with fresh water (Aharon Oren, personal communication). A third example is the deep anoxic basins in the Mediterranean, where the water is nearly saturated with $MgCl_2$ (5.0 *M*, a_w ~0.3) (van der Wielen *et al.*, 2005). The presence of microbes in this brine is indicated by 16S ribosomal RNA genes and some enzymatic activity. However, there is no direct evidence of reproduction or growth in the brine—the DNA and enzymes could ultimately be derived from microbes that grew in overlying water with much lower salinity rather than in the highly concentrated brine.

The lowest solute-induced a_w for which well-documented growth has been shown is 0.62. This is the case of xerophilic fungi growing in highly concentrated (83% wt/vol) sucrose solutions (Harris, 1981). Sucrose solutions as microbial habitats are more relevant to food microbiology than to naturally occurring environments such as brines or soils. Nonetheless, this value of a_w serves as a useful benchmark.

The lowest matric-induced a_w that allows microbial proliferation is dictated by solute diffusion and the availability of nutrients in solution. The lowest matric-induced a_w enabling growth of bacteria in culture is approximately 0.88. More importantly, the a_w at which microbial respiration becomes negligible as a soil loses moisture is approximately 0.86–0.93 (Sommers *et al.*, 1981). Soil respiration is a culture-independent measure and, thus, serves as a good indicator of the metabolic capabilities of all soil microbes. The actual a_w at which microbial proliferation ceases is, in all likelihood, higher than this in that soil microbes can respire by endogenous metabolism under conditions that are too dry for cell proliferation.

Water in contact with ice deserves special attention. The a_w of pure liquid water at any temperature is 1.0 and is not temperature-dependent. However, the a_w of ice is temperature-dependent and declines from 1.0 as temperature decreases. The a_w of ice is equal to the water vapor pressure over ice divided by the water pressure over pure liquid water. Thus, at $T = 0°C$, a_w of ice = 1.0; at $T = -20°C$, $a_w = 0.82$; at $T = -40°C$, $a_w = 0.67$; and so forth. Note that relative humidity meters (*e.g.*, Vaisala humicap sensors) read aw, and so a relative humidity meter placed in an atmosphere in equilibrium over pure ice at −40°C will read 67%.

The a_w of any solution in equilibrium with ice will be equal to the a_w of the ice and does not depend on which molecules are in solution or their quantity (Koop, 2002). Physically, the solution will gain or lose water until the a_w is equal between the solid phase (ice) and the liquid phase (the solution). This allows the a_w of ice-rich regions on Mars to be predicted solely from a measurement of temperature. Similarly, the eutectic temperature of any solution can be predicted since that is the temperature at which the a_w of ice is equal to the a_w of the saturated solution.

<u>FINDING</u>. Based on current knowledge, terrestrial organisms are not known to be able to reproduce at an a_w below 0.62; with margin, an activity threshold of 0.5 is proposed for use when considering special regions.

5D. Other Possible Limits to Terrestrial Life

SR-SAG concluded that a number of factors (some listed in Table 1, some not) contribute to a reduction in the probability of propagation, but for none except temperature and water activity is it possible at the present time to define practical threshold criteria that would apply to all terrestrial microbes.

The nutritional requirements for terrestrial microorganisms on Mars were considered to be key factors in limiting the proliferation of microorganisms on Mars. Terrestrial microorganisms require exogenous sources of nutrients, and accessible organic and/or inorganic nutrients in martian regolith have not been demonstrated (Biemann *et al*., 1977; Biemann and Lavoie, 1979). Although terrestrial chemoautotrophs do not require organic nutrients, they do require exogenous nutrient and energy sources, not all of which can be obtained in gaseous form. The diurnal temperature fluctuations shorten durations at temperatures above the minimum required for growth and require organisms to be capable of surviving repeated exposure to eutectic freezing. Both elicit a stress response that diverts resources toward repair of cell damage rather than cell division. The strong biocidal UVC irradiation on Mars helps to further constrain the proliferation of terrestrial microorganisms on Mars by two key processes: (a) UVC irradiation can quickly reduce the viability of sun-exposed bioloads on spacecraft surfaces, and (b) UVC irradiation will likely reduce long-distance dispersal of the remaining viable bioloads by imposing a highly lethal nonionizing radiation environment on the dispersed microorganisms. [The UV irradiation on Mars is significantly higher in the UVC region (190–280 nm) than on Earth because of a generally thinner atmosphere and the lack of an extensive ozone layer (Kuhn and Atreya, 1979; Appelbaum and Flood, 1990; Cockell *et al*., 2000; Patel *et al*., 2002). On Earth, the ozone layer attenuates all UV irradiation below 290–300 nm, *i.e.*, no UVC wavelengths reach the Earth's surface. The presence of UVC irradiation on Mars creates an environment at the surface that exhibits a total UV flux (200–400 nm) that is up to three orders of magnitude more biocidal than on Earth (Cockell *et al*., 2000; Patel *et al*., 2002). Recent models suggest that the high UVC flux on Mars can act to reduce the viability of some sun-exposed microbial cells on spacecraft surfaces by greater than six orders of magnitude in as short a time as a few tens of minutes to no more than several hours (Schuerger *et al*., 2003, 2006; Newcombe *et al*., 2005). In addition, the downwelling UVC will penetrate pits, cracks, and other microscopic topographical features on spacecraft materials, resulting in some of the more sheltered microorganisms becoming inactive in reasonably short periods of time (Schuerger *et al*., 2005). However, the biocidal effects of UVC cannot reach deeply embedded bioloads, cannot penetrate UV-absorbing materials, and cannot affect bioloads on internal components of spacecraft.] For organisms near or at the surface, long-term exposure to galactic cosmic rays and solar particle events will certainly increase lethality and reduce viability.

None of these secondary factors has been adequately measured or modeled for the martian surface or near-subsurface to allow us to set thresholds about their effect on survival, growth, and proliferation of microorganisms on Mars. However, all combine to lower the likelihood that Earth organisms will be able to propagate or even spread at the surface while remaining viable.

<u>FINDING</u>. **Despite knowledge that UV irradiation at the surface of Mars is significantly higher than on Earth, UV effects have not been adequately modeled for the martian surface or near-subsurface to allow us to set thresholds about their effects on growth and proliferation of microorganisms on Mars. However, UV may be considered as a factor that limits the spread of viable Earth organisms.**

5E. Discussion

We conclude that thresholds for temperature (−20°C) and a_w (0.5) define conditions below which Earth organisms will not grow or replicate. Such conditions that might exist on Mars must actually exceed both of these parameters for periods of time sufficient to allow growth and cell division to occur. We consider these to be very conservative values. Cell division has never been observed below sustained temperatures of −12°C, and 0.5 a_w is much lower than the minimum value for matric-induced a_w values that allow for microbial propagation in terrestrial environments. This value is more conservative (lower) than the lowest solute-induced a_w known to be compatible with growth: the unusual case of yeasts growing in a concentrated solution of sugar. Modeling studies predict that long-term conditions exceeding these thresholds will not persist long enough to permit cell division cycles, which may require weeks to years for completion.

Although it is impossible to assign with certainty values for probability of growth of an Earth organism on Mars, we can be confident that assignment of "special region" requires that conditions exceed minimal temperature and a_w parameters defined above. In addition, the litany of environmental stressors discussed above further reduces the likelihood of propagation of terrestrial organisms.

FINDING. The most practically useful limits on the reproduction of terrestrial microorganisms are temperature and a_w, for which threshold values (with margin) can be set at −20°C and 0.5, respectively.

6. WATER ON MODERN MARS

Water on Mars is best analyzed in two broad, distinct classifications: the parts of Mars that are at or close to thermodynamic equilibrium and those that are in long-term disequilibrium.

6A. The Distribution of Water Where It Is at Equilibrium

Introduction. Numerical thermodynamic models of martian surface and subsurface temperatures have been successfully used for decades to examine the physical nature of the surface layer (*e.g.*, Kieffer *et al.*, 1977) and the behavior of subsurface volatiles (*e.g.*, Leighton and Murray, 1966). The repeatability of thermal inertia results from data set to data set (*e.g.*, Jakosky *et al.*, 2000) indicates that these models are generally accurate to better than a few degrees during most seasons and even more accurate on an annual average.

The *absolute* humidity (*i.e.*, the partial pressure of water) varies with time and location on Mars, but it seldom climbs much above 0.8 μbar. *Relative* humidity is the ratio of this partial pressure to the saturation vapor pressure of the air or regolith, which is a function of temperature, varying exponentially with $1/T$. Over the large temperature extremes of a martian day, the relative humidity may go to 100% at night as frost is deposited and fall to very low values in the warmth of the day, but the absolute humidity will vary very little. Where ice is in equilibrium with the observed atmospheric water vapor pressure on modern Mars (*i.e.*, when it is at the frost point), it will have a temperature of about −75°C (Mellon *et al*., 2004). This means that, where there is vapor diffusive equilibrium with the atmosphere, ice is unstable with respect to sublimation at temperatures above −75°C, and water vapor is unstable with respect to freezing at temperatures below that.

Mars is warmer at the equator than at the poles. Using factors like the thermal inertia of the surface material and the solar insolation (which may include slope effects), it is possible to quantify this and develop planetary-scale maps of parameters like the fraction of the upper meter that is composed of ice and the depth to the ice table (*e.g.*, Chamberlain and Boynton, 2006; Mellon and Feldman, 2006; Aharonson and Schorghofer, 2007). Such models (*e.g.*, Fig. 3) have a general structure that consists of abundant ice within 1 m of the surface at high latitude, a mid-latitude belt of ice at a depth of 1–5 m, and an equatorial belt where ice is either deeper than 5 m or absent altogether. The steady-state ice depth depends on thermal properties and is independent of molecular diffusivity.

Equilibrium thermodynamic models show that the depth to the top of the ice table increases abruptly at about 50° latitude in both the north and south hemispheres (Fig. 4). This has been studied extensively (*e.g.*, Farmer and Doms, 1979; Paige, 1992). It is typical in model results for the transition from a depth of 5 m to infinite to occur in less than a degree of latitude. Thus, in these kinds of models there is no practical distinction between ice table depths of 3–10 m, which is the maximum depth of penetration for crashes that involve currently envisioned martian spacecraft.

A critically important value of such thermodynamic models is that they have predictive value down to spatial scales much finer than that achievable by observational data. At equilibrium, intensive variables like temperature and a_w are equal at all scales—this is one way to define equilibrium. This is essential to interpreting special regions, since the scale of spacecraft observation

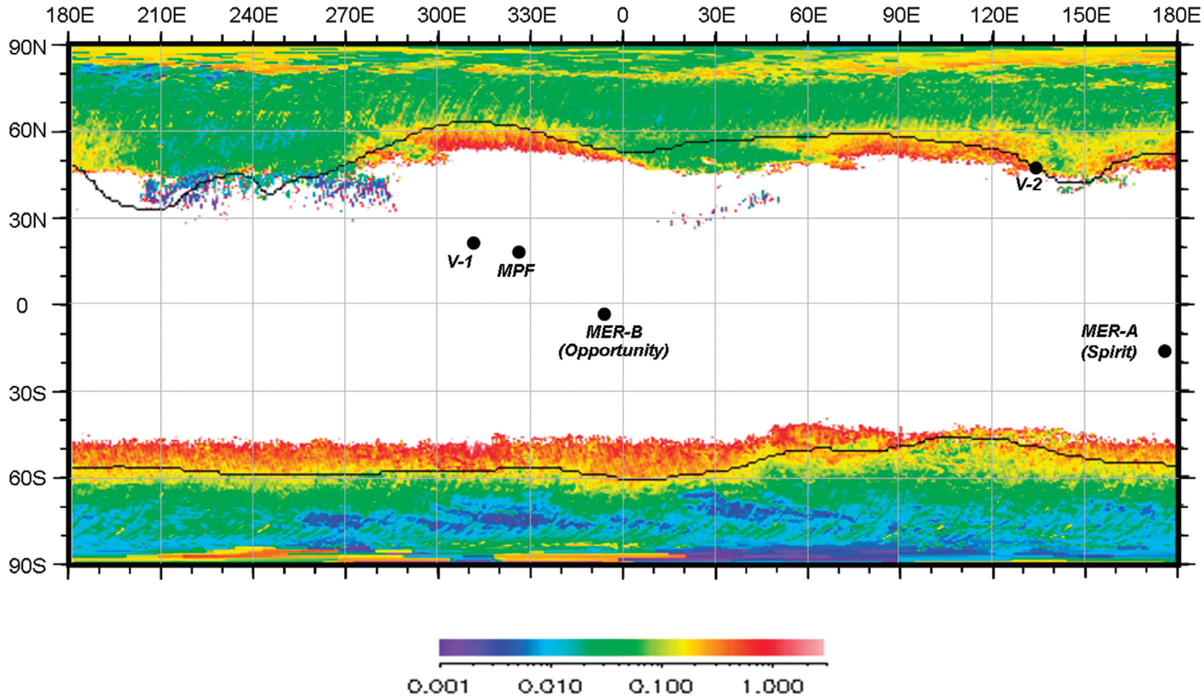


FIG. 3. Map of depth to the ice table (depth scale in meters), from Mellon and Feldman (2006), calculated assuming 20 precipitable microns of atmospheric water vapor scaled by elevation. This depth represents a 100–1,000 year average. The solid line is the 6 counts/s isopleth for epithermal neutrons (see Fig. 5). MPF, Mars Pathfinder; V-1, Viking-1; V-2, Viking-2.

(the footprint of the GRS instrument, for example, is approximately 3×10^5 km^2) can be many orders of magnitude larger than the finest scale of relevance for biology (microns). Note that the degree to which any martian environment does or does not approach equilibrium does not depend on whether ice is actually present; a_w is a property of both gaseous and solid phases. Similarly, the magnitude of heterogeneity in T and a_w depends on the effect and scale of geologic processes that can produce departures from equilibrium conditions (see Sec. 7 of this report). From our understanding of the Earth, it is known that there are macroscopic processes that can produce distinct departures from equilibrium, but the scale tends to be local to regional, not microscopic (for example, one grain in a rock is not at a meaningfully different temperature than the next grain).

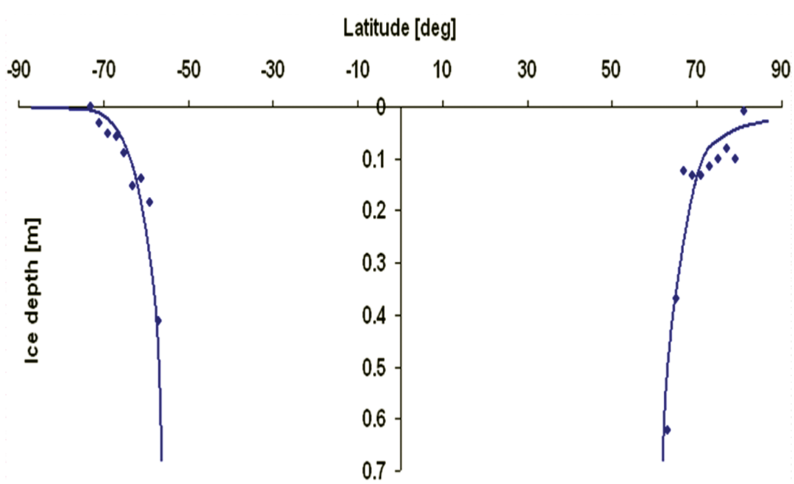


FIG. 4. Two cross-sectional profiles showing the depth to the ice table [presented by Hock and Paige (2006) at the Mars Water Conference]. Calculations are done for a longitude of 120°E in north and 220°E in south.

Is an equilibrium model consistent with observed data? The strong general agreement between models of ground temperature and ground ice and observations of temperature and hydrogen suggests that such numerical simulations capture the major portion of the relevant physical processes that control these phenomena. These models are based on well-known physical processes of solar heat, radiation, conduction, etc. They have been validated by analytic solutions and by the general consistency with spacecraft observations (including planets other than Mars). The errors in these models tend to be related to missing or oversimplified secondary physics. For example, emissivity variations from one region to another due to changes in mineralogy can affect the kinetic surface temperature and are usually not included in numerical simulations. The magnitude of these errors can be as much as a few degrees.

Comparison between Mars Odyssey GRS measurements (indicating the presence of subsurface hydrogen and subsurface ice) (Fig. 5) and theoretical models of ice stability based on these same thermodynamic numerical models demonstrates excellent agreement between theory and observation (Mellon *et al.*, 2004).

6B. Possible Secondary Factors That Affect a General Thermodynamic Model

6B-i. The Possible Effect of Diurnal and Seasonal Heating/Cooling

The martian surface is subject to diurnal and seasonal heating and cooling, which can cause significant temperature variation. These temperature fluctuations are attenuated in the shallow subsurface. As shown in Fig. 6, the scale of this attenuation depends on the thermal inertia of the surficial material. When no subsurface ice is present (*e.g.*, Fig. 6a), subsurface heating/cooling beyond a few degrees occurs only in the upper 2 m or so. However, when a subsurface layer of ice is present (Fig. 6b), it has the effect of wicking away the heat—the high thermal conductivity of ice resists the further propagation of the thermal wave, and significant heating can be restricted to much shallower depths (0.5 m in this example). Al-

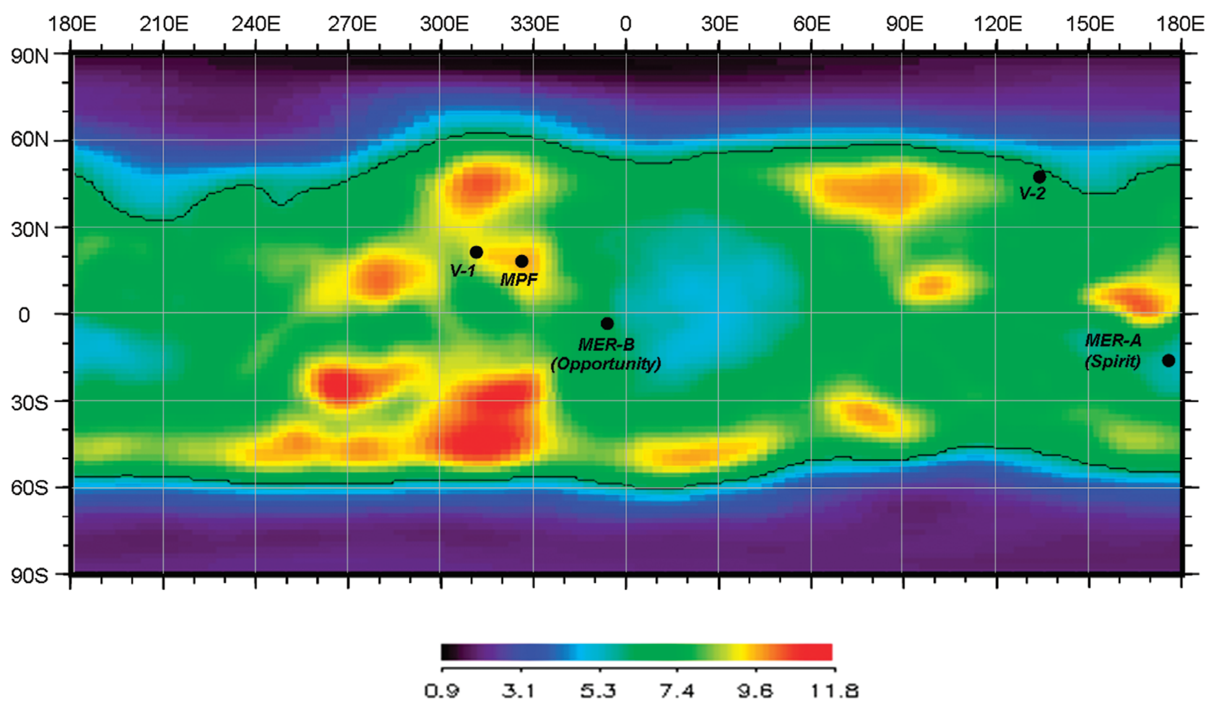


FIG. 5. Map of epithermal neutrons, which are very sensitive to subsurface hydrogen and water ice, from the GRS instrument on Mars Odyssey (Mellon and Feldman, 2006). Only summer data from both hemispheres are used (winter CO_2 frost obscures the ice signature by adding hydrogen poor mass atop the soil—seasonal CO_2 can be as much as a meter or more at high latitudes). Beyond a threshold boundary of 6 counts/s, ice detection falls off rapidly toward the equator. This boundary is more diffuse in the northern hemisphere than in the southern hemisphere. MPF, Mars Pathfinder; V-1, Viking-1; V-2, Viking-2.

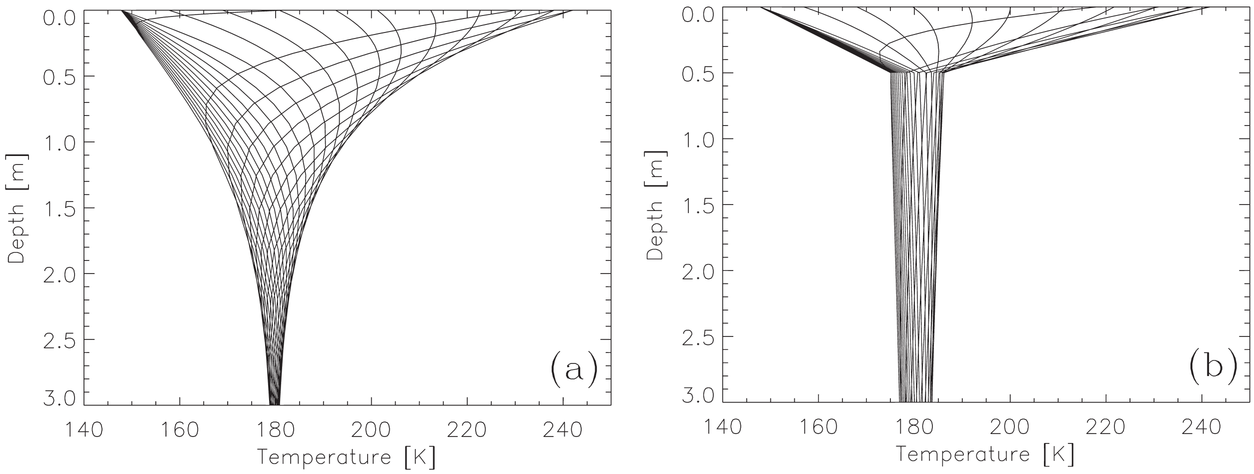


FIG. 6. Example subsurface temperature profiles for (a) a homogeneous subsurface and (b) a layered subsurface, from Mellon *et al.* (2004). Each curve is a diurnal average temperature profile superimposed at 25-day intervals for a full martian year. Both cases are for 55°S latitude for a thermal inertia of 250 J m^{-2} K^{-1} $s^{-1/2}$ and an albedo of 0.25. For the layered case the thermal inertia is increased at and below 50 cm to 2,290 J m^{-2} K^{-1} $s^{-1/2}$ to correspond to densely ice-cemented soil. The magnitude of the temperature oscillation is reduced by almost a factor of 5 at or below the ice table.

though Mars has an ample supply of near-surface water, it is stubbornly sequestered in solid form at temperatures below the frost point, either on the polar caps or in vast high latitude, subsurface locations (Leighton and Murray, 1966).

The surface of Mars at many low-latitude locations may exceed 0°C in the peak of the day, an observation that has been offered as possibly enabling the presence of liquid water. However, as discussed above, given the extremely low vapor pressure of water in the martian atmosphere, this temperature is 75°C above the frost point. Therefore, it would be impossible for new water to condense, and any previously present ice or water would quickly sublime or evaporate. Once in the vapor phase at these elevated temperatures, water in the shallow subsurface would tend to diffuse upward to the atmosphere or downward to a colder place. The thermal minimum in the subsurface would function as a cold trap. Cyclical heating and cooling of the uppermost martian crust would, therefore, result in progressive desiccation. Maps of locations that receive the most heating (Fig. 7) are equivalently the places that have been the most desiccated. In addition, it is worth noting that cyclical diurnal and seasonal warming causes rapid sublimation, while a cold fluctuation brings only *slow* ice accumulation, simply because the atmosphere does not supply a significant source of water. [By analogy (for those of us old enough to remember) it takes only a short time to defrost a freezer, but a relatively long time for the ice to accumulate again.]

Even though the temperature maxima may exceed 0°C at the surface, it is possible to show, from a map of the mean surface temperature (*e.g.*, Mellon *et al.*, 2004) and the general shape of the temperature attenuation curves (Fig. 6), that the temperature 10–20 cm below those surfaces remains perpetually below −40°C.

6B-ii. The Possible Effect of Recharge from Subsurface Water Reservoirs

As discussed above, at localities where the regolith is permeable to gas (which is certainly the case for most or all of Mars), there will be vapor-diffusive exchange between the atmosphere and ice within this volume. This exchange involves two-way mass transfer from ice into vapor, and from vapor into ice. This process leads to the formation of an ice table, where there can be a high concentration of ice below the equilibrium point and none above it. This is a stable condition, and one that can last indefinitely. As discussed by Clifford (1991, 1993), near-surface ground ice can also be replenished by reservoirs of H_2O in the deeper subsurface. The existence of deep reservoirs of H_2O at equatorial latitudes on Mars has been postulated by a variety of authors based on a variety of arguments. Water vapor from such

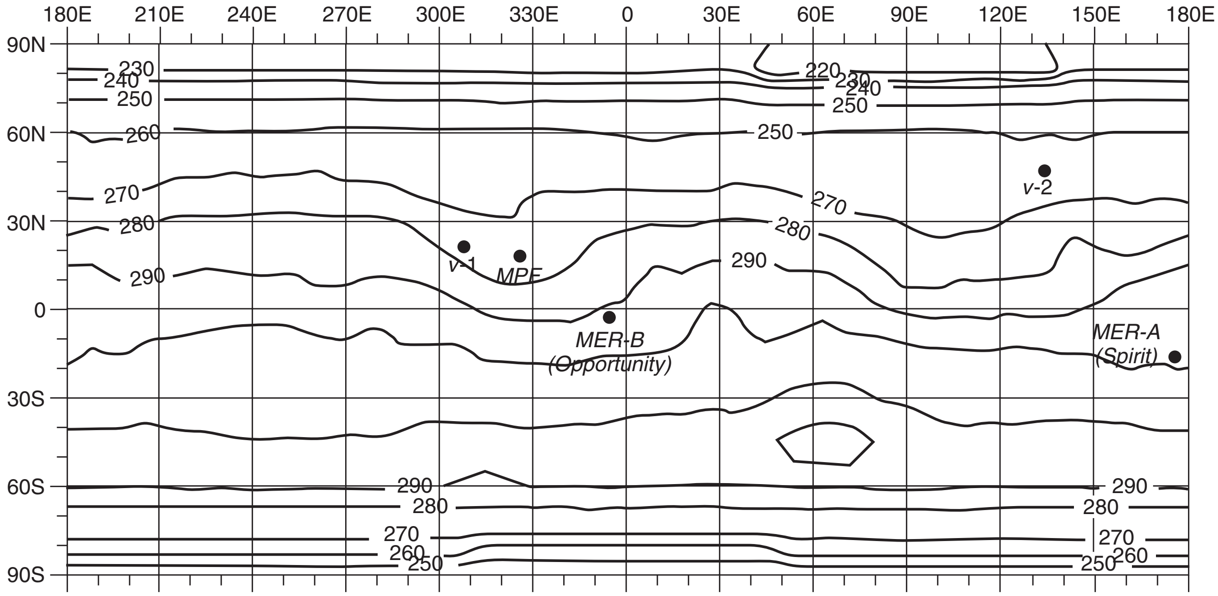


FIG. 7. Peak surface temperature on Mars (from Haberle *et al.*, 2001). The warm areas correspond to the most arid spots. MPF, Mars Pathfinder; V-1, Viking-1; V-2, Viking-2.

reservoirs could migrate up the geothermal gradient to the thermal minimum in the shallow subsurface, and from there sublime into the atmosphere.

The presence or absence of ice in the shallow martian subsurface depends primarily on the stability of ice. Subsurface vapor plumes will stay in a vapor form unless the temperature is below either the frost point or the dew point—above that, neither ice nor water will form. It does not matter whether vapor is being contributed from one source or two. In brief, the ice exists where it exists because it is cold, even when it is replenished. If any location warms to the point that it approaches the biological threshold (−20°C or more), then the thermal gradient will work in the opposite direction and drive water up and down. Seasonal variation does not change this conclusion. Diffusion will occur rapidly under summer conditions (warm to cold) as compared with wintertime (cold to colder), and therefore, the dominant direction of flow will always be out of the thermally fluctuating zone. (This is why the surface layer stays dry in the subpolar regions on Earth.)

A case where subsurface recharge might matter to the analysis of the upper 5 m is when a surfacial crust of very low permeability is present and the rate of recharge from below significantly exceeds the rate of diffusive loss to the atmosphere. This could cause the partial pressure of water in the shallow subsurface to go up, which would in turn cause the frost point to increase. This situation is discussed in detail in Sec. 6B-iv of this report. In summary, however, we have no evidence that such permeability barriers exist on Mars, and arguments can be developed for why they are geologically implausible.

6B-iii. The Possible Effect of Unfrozen Thin Films of Water

Since there is known to be water in the martian atmosphere (about 8 μbars), as well as water cycling at some rate between crustal and atmospheric reservoirs of water, it is inevitable that thin films of water are present on mineral grains in the dry parts of the martian crust. The "stickiness" of water is well known to experimentalists who operate high-vacuum equipment on Earth. The bad news is that there is no way to make a direct measurement of the thickness of thin films of water in different martian environments. However, the good news is that the a_w in the thin films, of whatever thickness, can be calculated from the relative humidity of the atmosphere in equilibrium with the thin film. As shown below, the aw, the temperature, or both are less than the biological thresholds across the entire martian surface and shallow subsurface.

6B-iv. The Possible Effect of Semipermeable Crusts

On Earth, soil crusts can provide a significant permeability barrier, through which the rate of fluid flow can be lower than the rates of resupply or fluid loss on either side of the barrier. In such cases, water can be trapped in a transient way, even when it is out of equilibrium with the atmosphere.

Observed crusts on Mars. Crusts are common at the martian landing sites visited through 2005. Observations to date show them to be relatively weak and friable. Viking "duricrusts" at Chryse Planitia (Viking-1) were readily broken by digging action, and those at Utopia Planitia (Viking-2) were disaggregated simply by shaking them in the acquisition scoop (Clark *et al*., 1982). Many crusted materials have been seen at both Mars Exploration Rover (MER) landing sites, but all seem to be easily broken as the wheels pass over them (Richter *et al*., 2004; L. Richter *et al*., manuscript in preparation). To date, no examples of high-strength crusts have been discovered at any of the five landing sites. Although we have no data on the permeability of any of these crusts due to their friability, discontinuous nature, and porosity, they do not appear to be particularly impermeable.

Terrestrial analogs. Since other types of crusts might exist elsewhere on Mars, some of which may be less permeable, it is important to consider other kinds of crusts known from terrestrial experience. Surface crusts associated with soils on Earth are classified as biological, chemical, or physical (Soil Survey Staff, 1999):

- *Biological crusts* are composed of mosaics of cyanobacteria, green algae, lichens, mosses, microfungi, and other bacteria (Belnap *et al*., 2001).
- *Chemical crusts* are largely formed where water containing dissolved salts, commonly carbonate, sulfate, and chlorides, accumulates in shallow depressions allowing evaporation and precipitation at the surface. Common settings for chemical crust include dry lakebeds or sabkas. Salt crust may also form at the soil surface from capillary rise of salt-rich soil moisture.
- *Physical crusts* primarily result from the formation of aggregates from a reconstituted, reaggregated, or reorganized layer of mineral particles. Common types include structural (*e.g*., raindrop impact), depositional (surface flooding), freeze-thaw, and vesicular. Aggregates can range from $\sim 10^{-2}$ to 10^2 mm in diameter, with the larger aggregates due to the formation of soil structure.
- Another type of soil crust is the strongly cemented subsoil layer where the soil matrix has been cemented by the extensive accumulation of carbonate, salt, and silica (*e.g*., duricrust, caliche).

Common attributes among all types of surface crust is that they generally enhance surface sealing, provide surface stability, limit wind and water erosion, increase aggregation of binding of soil particles, and are commonly $<$10 cm thick.

The best terrestrial crust-forming analog for the martian surface is, perhaps, a type of physical crust referred to as vesicular crust, which is common to desert regions on Earth. This is associated with reg soils or desert pavements, features ubiquitous to nearly all arid deserts (McFadden *et al*., 1998). Vesicular crusts typically underlie a single surface layer of cobbles or gravel. Desert crusts are primarily derived from the long-term accumulation of aeolian dust (particle diameters $<$0.1 mm) and require 10^3–10^5 years to form (McDonald *et al*., 1995). The density of this type of crust ranges from about 1.5 to 1.9 g/cm^3, and they are commonly 3–10 cm thick (McDonald, 1994).

Permeability of desert crust. The measured saturated hydraulic conductivity of desert crust typically ranges from 0.75 to 0.5 cm/h. Actual conductivity is typically lower than that measured at saturated conditions because of trapped air within the crust (McDonald *et al*., 1996; McDonald, 2002; Young *et al*., 2004; Meadows *et al*., 2005). Terrestrial desert crusts are not completely impermeable because a wide range of processes, primarily dispersive stress and tensional release, result in the formation of voids, pores, and fractures that prevent the continuous sealing of the soil matrix. Even the most cemented layers (*e.g*., caliche, duricrust) always have fractures that limit the horizontal and vertical extent of cementation and sealing. Although the formation of crust fractures is exacerbated by biological processes (*e.g*., root propagation), dispersive or tension stress is first required to promote development of fractures. On the surface of Mars, processes such as freeze-thaw, formation of ice, and ground shaking due to seismic activity and meteorite impact are likely to enhance formation and propagation of fractures in crusted materials.

Episodic retention of moisture beneath desert crust on Earth has been observed to happen when the rate of water recharge (on Earth, primarily via rain) exceeds the rate of water loss (*i.e.*, vapor loss through the crust). For example, measurements in soil beneath physical crusts in hyperarid deserts show that the crust enhances moisture retention and leads to a lower soil temperature (~3–6°C) relative to soils that lack physical crusts, typically for a period of several months (E.V. McDonald, personal communication, 2006). However, such anomalies are dynamic and typically decay on a time scale of a year.

For such cases, the mean rate of diffusive gain will, in time, equal the mean rate of diffusive loss, and equilibrium will be attained (although both rates will be lower than when crust is absent). Although other crust-related processes may be discovered in the future, SR-SAG concludes that, to within its standard of confidence, this is not an environment that will lead to the presence of liquid water.

FINDING. **Although soil crusts on Mars have been observed at the past landing sites, and other crust types are hypothetically possible elsewhere, experience with desert crusts on Earth shows that the effect of a semipermeable crust is to retard, not prevent, the achievement of equilibrium.**

6C. Calculation of Water Activity on Modern Mars

Persistent liquid water at or near the martian surface thus requires a significant departure from the general planetary setting in the form of either long-term disequilibria (caused either by geothermal sources of heat and water or by vestigial sources from prior climates that have survived for 10^4–10^7 years by virtue of giving up their water *very* slowly) or short-term disequilibria (from the influence of a spacecraft on a cold, icy site or a transient event such as a meteorite impact). If any surface or shallow subsurface location on Mars were to warm to the biological threshold (−20°C or more), then the heating would drive water up (to the atmosphere) and down (to a colder place). Seasonal variation does not change this conclusion. Diffusion will occur rapidly under summer conditions (warm to cold) as compared with wintertime (cold to colder), and, therefore, the dominant direction of flow will always be out of the thermally fluctuating zone. This is the reason why, on Earth, we observe that the surface layer stays dry in the subpolar regions.

On Earth, life can exist and propagate in soils we might casually consider "dry," often surviving on thin films of water in capillaries or at grain boundaries. The a_w was introduced in the previous section as a quantitative measure of dryness, and an a_w threshold was established below which terrestrial life is not known to be able to reproduce. The a_w, a measure of availability of water, is defined as the relative humidity in the pores of the soil (although expressed as a decimal fraction rather than a percentage). At a microscopic level, it is typically surface tension associated with the concave geometry of water or ice droplets that holds vapor pressure below nominal saturation, which results in a_w values less than unity. Alternatively, salt content can lower the saturation vapor pressure as well as the melting point. The a_w is thus a proxy for the specific physics and bioavailability of thin films of water needed for microbe propagation, and obviates the need to consider specific soil properties or the presence of brines. Moreover, if soil is in equilibrium with the surrounding atmosphere, then a_w can be determined directly from the atmospheric relative humidity. In a single, easily determined parameter, we can capture the detailed microscopic interactions of microbes, water, and soil.

Figure 8 shows the equilibrium a_w of martian soil as a function of temperature, derived by assuming the absolute humidity to be 0.8 μbar, in equilibrium with the atmosphere. In warm soil, a_w is literally orders of magnitude too small to support life—there simply is not enough water to dampen the soil sufficiently. The a_w approaches unity at the frost point, but at a temperature far too low to support life. The box in the upper corner of Fig. 8 delineates the conditions under which terrestrial life could propagate, far from the water equilibrium.

FINDING. **Where the surface and shallow subsurface of Mars are at or close to thermodynamic equilibrium with the atmosphere (using time-averaged, rather than instantaneous, equilibrium), temperature and a_w in the martian shallow subsurface are considerably below the threshold conditions for propagation of terrestrial life. The effects of thin films and solute freezing point depression are included within the a_w.**

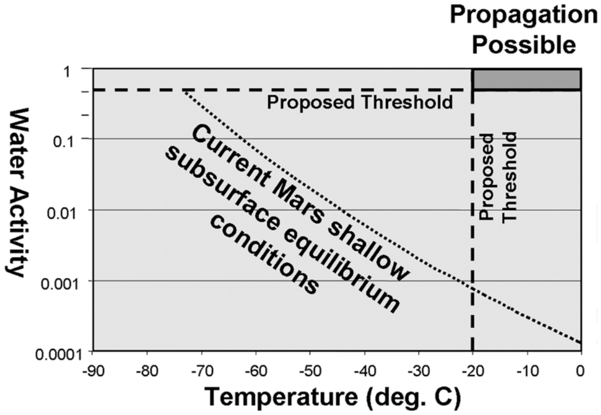


FIG. 8. The a_w of present-day Mars versus temperature in equilibrium with the present-day atmosphere, assumed to have a water partial pressure of 0.8 μbar. The region of concern for life propagation is shown in the upper right. The a_w of pure ice is less than 1 over the range of temperatures on Mars (*e.g.*, 0.46 at -90°C). The a_w of materials on Mars at subfreezing temperatures is always less than the value of the a_w of pure ice. This is a small correction and is not shown here.

7. MARS ENVIRONMENTS IN THERMODYNAMIC DISEQUILIBRIUM

7A. Introduction

Section 6 of this report argued that, where Mars is at or close to long-term thermodynamic equilibrium, the threshold conditions for propagation of terrestrial organisms are not met anywhere at the martian surface and shallow subsurface. However, there remains the possibility that some parts of Mars are not at equilibrium (Carr, 1996). For the purpose of this analysis, we distinguish short-term disequilibrium (*i.e.*, the changes in heating that occur on a daily or annual cycle) and long-term disequilibrium (changes that happen as a result of geologic processes with a time constant longer than 1 year). Long-term disequilibrium conditions are the subject of this section.

All over the planet, the daily and annual temperature cycles result in heating and cooling within the outermost skin of Mars. Within this process, the positive and negative excursions from equilibrium offset each other—on average, the material is at equilibrium. This has been described as a "dynamic equilibrium." In such an environment, any liquid that might form at the higher temperature would be transported in a matter of *hours* to one of the cold ice reservoirs by the process of evaporation and condensation. This would have the effect of leaving the surface perpetually desiccated. For the purpose of identifying and evaluating environments in long-term disequilibrium, some workers (*e.g.*, Hock and Paige, 2006) have set up an "annual equilibrium" criterion, and then look for excursions from that.

One of the implications of the laws of thermodynamics is that systems tend to move toward equilibrium. Thus, an environment in long-term disequilibrium is one where water and temperature were in equilibrium under conditions at an earlier time, but those conditions have changed and do not hold for the present. Geological deposits formed under such conditions will seek the modern equilibrium. As discussed below, there are several examples where the path *could* take such deposits through a liquid water field as they adjust to new conditions. Long-term disequilibrium environments might survive for 10^4–10^7 years by virtue of giving up their water very slowly.

Long-term disequilibrium develops in response to certain geological processes. These processes operate at different rates and at different times and exhibit different kinds of geologic and geomorphic manifestations. In evaluating martian environments, there are two ways to proceed, both of which are used in this report:

- *Description of processes*. These constitute "theory" for integrating a series of observations (for example, geothermal vent). In some cases, there may be multiple working hypotheses to explain a given observation. There are also processes that are currently hypothetical because the predicted observation has not yet been recorded (but in some cases is the focus of active search).
- *Description of geomorphologic features*. The acquisition and analysis of orbital images generate these kinds of basic data, but the linkage of the observations to the inferred geological processes is interpretive and further aided by supporting data from other approaches (for example, dissected mid-latitude mantled terrain).

7B. Gullies

Martian middle- and high-latitude gullies are geomorphic features whose age and origins are not fully understood, and there is a very real possibility that their genesis involved liquid water. The geomorphology and stratigraphic relations of these landforms to adjacent features suggest that some might be so young that they could be sites at which liquid water can occur, at least for brief periods, on the martian surface today.

Description. Mid- and high-latitude gullies were first described by Malin and Edgett (2000). The largest examples can be seen in Mars Odyssey Thermal Emission Imaging System (THEMIS) and Mars Express High Resolution Stereo Camera (HRSC) (an instrument on the 2003 Mars Express spacecraft) images, but the vast majority of these landforms are small enough that they are best recognized and described using images of better than 7 m/pixel, such as those from the 1996 MGS Mars Orbiter Camera (MOC) (Fig. 9).

Gullies always have a channel and usually exhibit an apron, unless it has been buried. Channels are commonly banked and, in some cases, meandered, but straight (not banked or meandered) examples also exist. Some gully channels are leveed. Gullies often (but not in all cases) exhibit an alcove above the channel; these form by undermining, collapse, and dry mass movement of debris (Malin and Edgett, 2000). Gully channels commonly originate at a point about 200–800 m below the local surface outside of the depression in which the feature occurs, and the alcove, if it is present, occurs above the point at which the channel begins (Malin and Edgett, 2000; Gilmore and Phillips, 2002; Heldmann and Mellon, 2004). Gully aprons, in some cases, are made up of dozens to hundreds of individual flow lobes. The majority of the tens of thousands of gullies identified in spacecraft images occur in the walls of craters, troughs, valleys, pits, and depressions. However, some variants on the theme are found on dune slip faces, crater central peaks, and the mountains surrounding Argyre Planitia (Malin and Edgett, 2000; Baker, 2001; Reiss and Jaumann, 2003).

Where Malin and Edgett (2000) lumped all of the gullies in this range of settings into a single group, Edgett *et al*. (2003) suggested that today they should be split into subgroups and that their differences may imply differences in how they form and whether a volatile is involved. The study of gullies is ongoing, and little research has yet addressed the geomorphic details that distinguish, for example, the gullies formed on dunes versus crater walls versus crater central peaks and the mountains rimming Argyre. Some gully-like forms occur at equatorial latitudes, but they do not exhibit all of the relevant morphologic criteria that distinguish the middle and high latitude landforms. Specifically, the equatorial features are (a) the straight, narrow avalanche chutes and attendant talus deposits that have formed on some steep slopes among light-toned, layered rock outcrops in the Valles Marineris and associated chaotic terrains and (b) the abundant alluvial fans that occur within a unique impact crater, provisionally named Mojave, located at 7.6°N, 33.1°W (Williams *et al*., 2004).

Planetary distribution. Nearly all cases occur at a location poleward of 30° latitude in both hemispheres (Fig. 10). Gullies that occur equatorward of 30° are very rare; the majority of these are poleward of 27° and mainly located on the north walls of Nirgal Vallis (Malin and Edgett, 2000; Edgett *et al*., 2003; Balme *et al*., 2006).

Possible relationship to water. The origin of the gullies has been much discussed and debated over the past 6 years, but no single explanation has yet satisfied all investigators interested in the subject. The majority of published results regarding mid- and high-latitude gullies have centered on the hypothesis that liquid water is involved and that the geomorphic expressions of the banked channels, tributary channels, meandering channels, and flow lobes in apron deposits are all clues regarding the rheologic properties of water-rich debris flows that have come through a given gully channel on more than one occasion (Malin and Edgett, 2000; Hartmann *et al*., 2003).

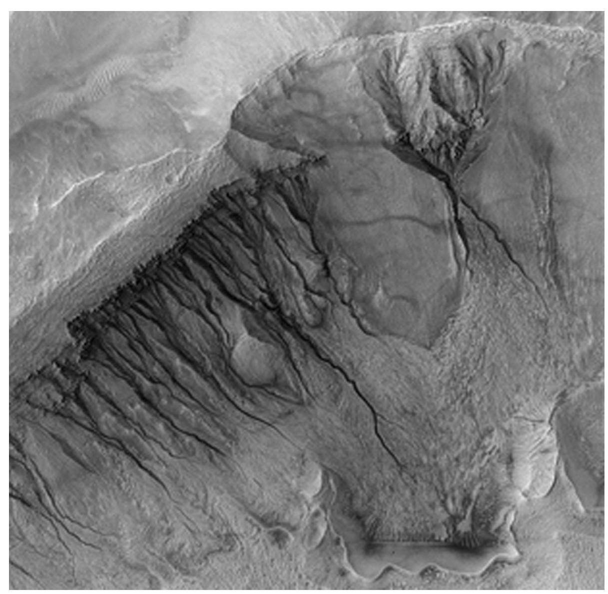


FIG. 9. Typical mid-latitude gullies on the wall of a crater located at 39.0°S, 193.9°E. This is a subframe of MOC image E11-04033; it covers an area ~3 km wide. The upper left corner of the image is the surface outside the crater; topography slopes down toward the lower right.

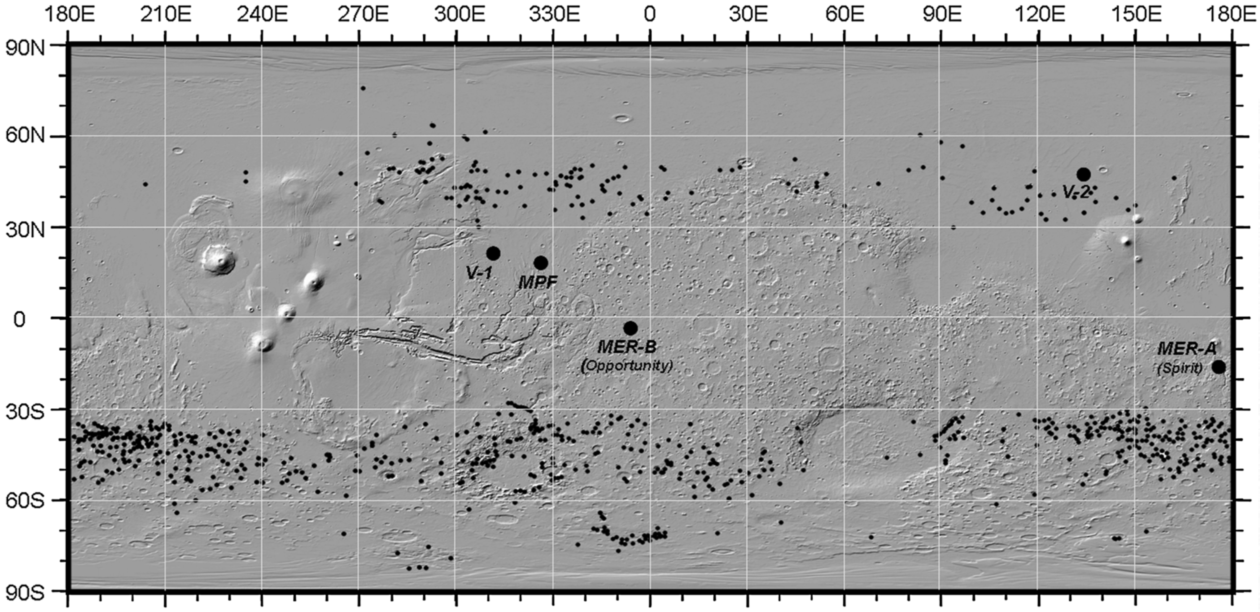


FIG. 10. Martian gully locations identified in MGS MOC images by K. Edgett and M.C. Malin through September 2005. Simple cylindrical projection; base is a shaded relief map derived from the 16 pixels/degree (or 3704.66 m/pixel) global MOLA DEM. The five sites on Mars where landed missions have returned scientific data are shown for reference: V-1 = 1975 Viking-1, V-2 = 1975 Viking -2, MPF = 1996 Mars Pathfinder, MER-A = 2003 Mars Exploration Rover 'Spirit', MER-B = 2003 Mars Exploration Rover 'Opportunity'.

Among the liquid water hypotheses are those that invoke groundwater (Malin and Edgett, 2000; Goldspiel and Squyres, 2000; Mellon and Phillips, 2001; Gilmore and Phillips, 2002; Heldmann and Mellon, 2004), melting of ground ice (Costard *et al*., 2002; Hecht, 2002), and melting of a surface-covering snow pack (Lee *et al*., 2001; Christensen, 2003). Most publications about gullies on dune slip faces also involve water as a factor in gully formation (Mangold, 2003; Mangold and Costard, 2003; Reiss and Jaumann, 2003; Miyamoto *et al*., 2004a), although it is acknowledged that CO_2 or dry granular flow may have contributed instead of or in addition to water. In cases where water is invoked to explain the origin of gullies, discussion has included the notion that pure water (Heldmann *et al*., 2005) or brines (Burt and Knauth, 2002) may have been the agent responsible for the observed landforms. The gullies occur in a wide range of settings, most of them very far from volcanic regions, which suggests that igneous-induced hydrothermal processes are not likely involved (Malin and Edgett, 2000).

Alternate hypotheses center on genesis by release of CO_2 that had been trapped below ground (Musselwhite *et al*., 2001; Hoffman, 2002), but Stewart and Nimmo (2002) concluded that it is very difficult to trap sufficient quantities of CO_2 beneath the surface. Others suggested that the gullies formed by dry, granular flow, which would not require the participation of a volatile such as water or CO_2 (Treiman, 2003; Shinbrot *et al*., 2004), but this hypothesis does not explain banked or leveed channels, aprons consisting of many flow lobes, or association of channel heads with specific rock layers.

Age. The age of the gullies is central to the concern as to whether these landforms represent "special regions." The critical issue is whether liquid water can come to the surface at the head of a gully channel and run down to and deposit material in the gully apron today or sometime during the next 100 years.

Estimates for the age of gullies are based on their general geomorphic and stratigraphic youth, and their lack of superimposed impact craters. Pulling these observations together, Malin and Edgett (2000) concluded that the gullies could be less than 1 million years old, but this estimate was based on virtually no information regarding the absolute age of any particular landform. Reiss *et al*. (2004) examined small impact craters superimposing aeolian megaripples in Nirgal Vallis, which are also superposed by gully aprons. In their research, Reiss *et al*. (2004) concluded that the gullies in Nirgal Vallis—assuming the approach to deriving absolute ages from impact crater size-frequency distributions provides quantities that approach the true age of features on Mars—must be younger than 3 million years and might even be younger than 300,000 years in age. Other investigators have focused on the role of obliquity excursions and whether gullies might only be active under different temperature and pressure conditions than exist today (*e.g.*, Costard *et al.*, 2002; Christensen, 2003; Bermann *et al.*, 2005). However, the work of Heldmann *et al.* (2005) argued that modern pressure and temperature conditions are a better fit when the measured run-out distances of gully channel/apron complexes are considered, which suggests that the gullies are episodically active today.

The small sampling of gullies available in MGS MOC images at the time Malin and Edgett (2000) published their results suggested that all gullies are quite young relative to the martian geologic time scale. Among that small sample, gully channels and aprons were seen to cut or superpose landforms that are otherwise considered to be relatively young, such as aeolian dunes, aeolian megaripples, and patterned ground similar to that found in terrestrial periglacial settings. Furthermore, some gullies have dark floors, which indicates that dust does not settle and persist on these surfaces for periods longer than a martian season. The absence of dust on dark channel floors might be attributed to aeolian redistribution on a sandy surface or to recent movement of material through the gully channel by non-aeolian processes (*e.g.*, runoff of a liquid).

MGS MOC has continued to collect new images of gullies, almost daily, since 2000, so the sample is much larger today. In the larger sample very good examples of gully aprons and channels that have been cut by fissures and faults, peppered with small impact craters, or superimposed by windblown sand have now been recognized (Edgett *et al.*, 2003). However, these "old" gully examples are in the minority, and in most cases where craters superpose the gully surfaces, multiple craters occur, which suggests that they are secondary to a larger impact that happened elsewhere. In other words, the number of craters on the landforms associated with a gully is not necessarily a good indicator of age.

The MOC team has recently (http://www.msss.com/mars_images/moc/2005/09/20/dunegullies/) described a case in which repeated imaging by MOC revealed the formation of a new gully on a sand dune slip face (Hellespontus region, west of the Hellas basin). The gullies formed sometime between 17 July 2002 and 27 April 2005 (Malin and Edgett, 2005). The setting of this gully (a sand dune field on the floor of a crater) is distinctly different from that of the mid-latitude gullies that form on crater walls, and strongly suggests that there is more than one gully-forming process. Further observation of Mars is clearly needed. The fact that the Hellespontus gully-forming event occurred during the current decade shows that it is possible that some material can flow through at least some gully systems in the modern era, but it also raises a flag that asks: is it possible for new gullies to form at a location where no gully was observed before?

There are no obvious geomorphic criteria that can be used to predict which gullies might become active in this century. MGS MOC is currently being used to monitor the hundreds of gully sites (tens of thousands of individual gullies) to look for additional evidence of change, including features that might indicate whether water is present and/or has flowed down the relevant slopes in recent years. Features being sought include new gullies, or channels within a preexisting gully complex, and gullies in which bright material—perhaps ice or salts—has appeared. Other than the fact that gullies are restricted to certain kinds of slopes, we do not have a means of predicting where a gully may form in a location where one did not previously exist. Thus, for the purpose of PP, concern should extend to gully-forming regions, not just to the specific preexisting features. The scale of these regions, however, is as yet undefined.

<u>FINDING</u>. **Some—although, certainly, not all—gullies and gully-forming regions might be sites at which liquid water comes to the surface within the next 100 years. At present, there are no known criteria by which a prediction can be made as to which—if any—of the tens of thousands of gullies on Mars could become active during this century, or whether a new gully might form where there isn't one today.**

7C. Mid-Latitude Geomorphic Features That May Indicate Deposits of Snow/Ice

The middle martian latitudes exhibit a variety of surficial geomorphic features that suggest to some investigators that ice-bearing materials were deposited over much of the surface at these latitudes, perhaps during a prior obliquity excursion. Although many questions remain, a small and slowly growing literature describes evidence that ice-rich materials may once have mantled mid-latitude terrain, covering intercrater plains, crater walls, and other landforms. Two features, in particular, are relevant here:

1. A ubiquitous mantle that was deposited and has since become roughened by erosion in geologically recent time. The texture of this mantle is latitude-dependent. For the purpose of this report, this deposit is referred to as the *mid-latitude mantle*.
2. Accumulations of materials most commonly found on poleward-facing slopes of mid-latitude topographic features, such as crater walls and massifs. For the purpose of this report, this deposit is referred to as *pasted-on mantle*.

Mid-latitude mantle. A layered deposit, estimated to be 1–10 m thick, mantles much of the surface of Mars between 30° and 60° latitude in the northern and southern hemispheres (Kreslavsky and Head, 2000, 2002; Mustard *et al*., 2001). The terrain once described as "softened" in Viking-era literature (*e.g.*, Squyres and Carr, 1986) is, at MGS MOC scale, actually "roughened terrain" (Malin and Edgett, 2001). Examples of the texture of the mantle are shown in Fig. 11.

A progression from smooth-surfaced mantle to roughened and pitted mantles has been observed and described briefly by Malin and Edgett (2001) and Mustard *et al*. (2001). In some places, the erosion reveals that the mantle is layered (Milliken and Mustard, 2003). Because of the latitudinal relationship, Kreslavsky and Head (2000) hypothesized that the smoothing was due to a climate-controlled deposition of ice and dust, and Mustard *et al*. (2001) proposed that the roughening resulted from sublimation of ice from a mixture of ice and dust that settled from the atmosphere to produce the mantle. As discussed earlier in this report, thermodynamic models indicate that in the mid-latitude regions shallow ice is unstable under current atmospheric conditions (Mellon and Jakosky, 1995; Mellon *et al*., 2004), potentially accounting for the desiccation interpreted to have occurred there (*e.g.*, Head *et al*., 2003a).

The mid-latitude mantle material exhibits several distinct morphologies that change as a function of latitude (Fig. 12). Lower latitudes of ~30–45° are characterized by regions of smooth, intact mantle adjacent to regions where the mantle has been completely stripped from the surface, whereas higher latitudes of ~45–55° commonly exhibit a knobby surface texture indicative of incomplete removal of the material (Milliken and Mustard, 2003; Milliken *et al*., 2003). Latitudes poleward of ~55° exhibit the least dissection and removal of the material, which suggests that the mantle deposit has experienced less erosion at these latitudes and may still be ice-rich beneath a thin layer of ice-free dust (Mustard *et al*., 2001; Milliken and Mustard, 2003). The Mars Odyssey GRS data also detected an increased abundance of H+ within the upper ~1 m of the surface at latitudes higher than ~55°, which supports the hypothesis that these regions might currently include near-surface water ice.

The latitude dependence of the mid-latitude mantle, its variations in morphology with latitude, and the symmetry between the northern and southern hemispheres suggest that the deposition and removal of this layer are related to global changes in climate (Mustard *et al*., 2001; Head *et al*., 2003a). These deposits do not appear to be forming today, but instead appear to have formed as a result of past geologic processes during earlier periods of higher obliquity. The mantle blankets all preexisting surfaces, independent of topography and surface composition, which suggests that it originated by airfall deposition of dust cemented by ice precipitated from the atmosphere during favorable climate conditions. Periods of high obliquity have been proposed to cause changes in the martian climate that result

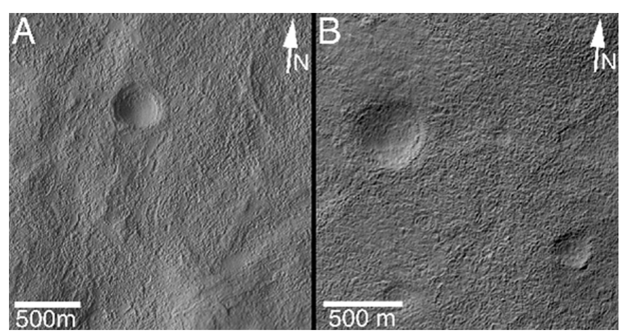


FIG. 11. Examples of mid-latitude roughened mantled surfaces in each hemisphere. (**A**) Northern hemisphere example from MOC image SP2-51906, near 30.0°N, 36.5°E. (**B**) Southern hemisphere example from MOC image M00-03091, near 40.1°S, 188.4°E. Both images are illuminated from the left.

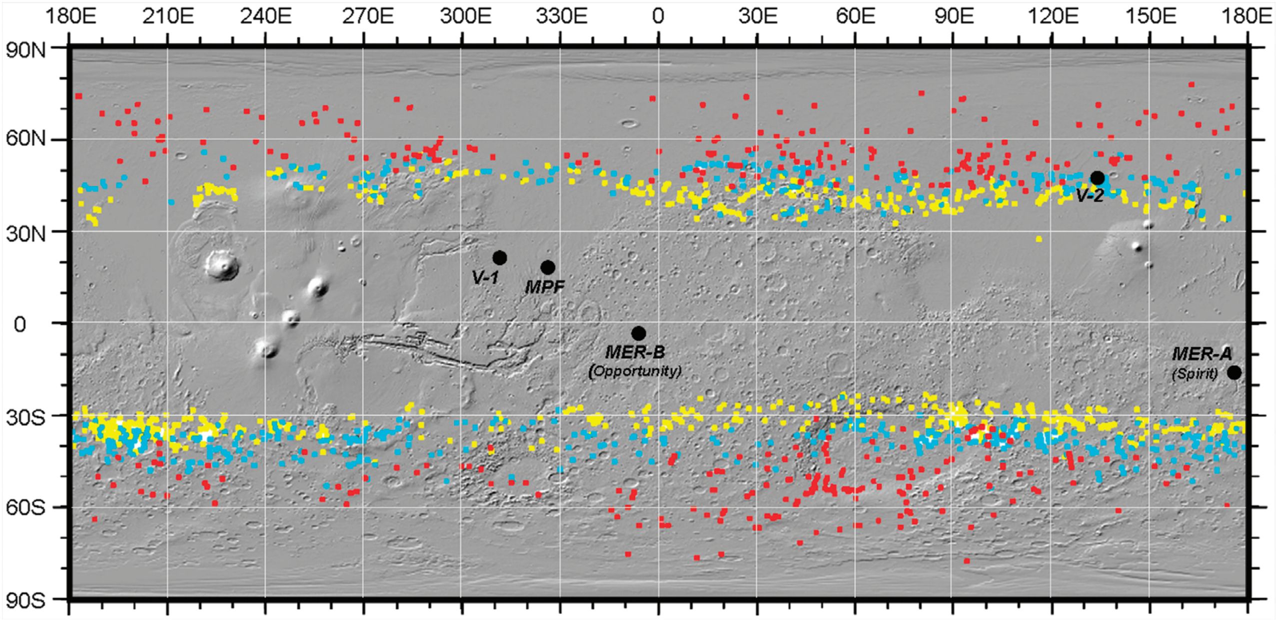


FIG. 12. Map showing the locations of MGS MOC images in which occur different erosion styles of the mid-latitude mantling materials (from Milliken and Mustard, 2003). Colors indicate: localized removal (yellow), knobby/wavy texture (cyan), and scalloped texture and total mantle cover (red). (For base map details, see legend to Fig. 10.)

in both an increase in atmospheric dust loading and a net transport of water from polar to mid-latitude regions. This provides a mechanism for multiple cycles of deposition and removal of ice-rich layers that is linked to orbital variations (Head *et al*., 2003a). The paucity of superposed craters, together with the correlations with recent periods of higher obliquity, suggested to Head *et al*. (2003a) that the mantle was emplaced between 2 and 0.5 million years ago, and has been undergoing sublimation and desiccation for the last half million years. In summary, these surficial ice deposits formed in an earlier geologic period, when Mars had a different pattern of surface insolation, and are no longer in equilibrium for the current orbital configuration. However, there is no evidence for melting over much of this region, and active layers are not predicted (Kreslavsky *et al*., 2006).

"Pasted-on" mantle. Among the wide variety of distinctive landforms found at middle latitudes, the one that is of concern here because of its apparently youthful age is that of mantles of material that appear to have been preferentially preserved on poleward-facing slopes in craters and on massifs and hills in both martian hemispheres (Fig. 13). Described colloquially among researchers in the Mars science community as "pasted-on" material, these mantling deposits were initially noted by Malin and Edgett (2001), who speculated that they bore some resemblance to accumulations of snow left behind on colder, more-frequently shadowed surfaces. However, the materials are not light-toned like snow. Mars Odyssey THEMIS visible images provided wider fields of view than MGS MOC, and thus greater vistas that show these "pasted-on" accumulations became readily apparent. This led Christensen (2003) to note that these accumulations seemed to occur most commonly on poleward-facing slopes at middle latitudes and to expand upon the speculation of Malin and Edgett (2001) that these might represent remnants of old snow accumulations.

Christensen (2003) and Milliken *et al*. (2003) proposed that the pasted-on mantle is a mixture of dust and ice (snow), and further proposed that they might be the source of water that creates mid-latitude gullies. However, the mutual relationship between gullies and the pasted-on terrain is not simple. There is a wide range of geomorphic attributes of mid-latitude slope-mantling materials and a wide range of mid-latitude gullies, and the relationships between them vary. For example, some gullies do not occur in association with such mantles. No one has yet published a detailed study on whether or how

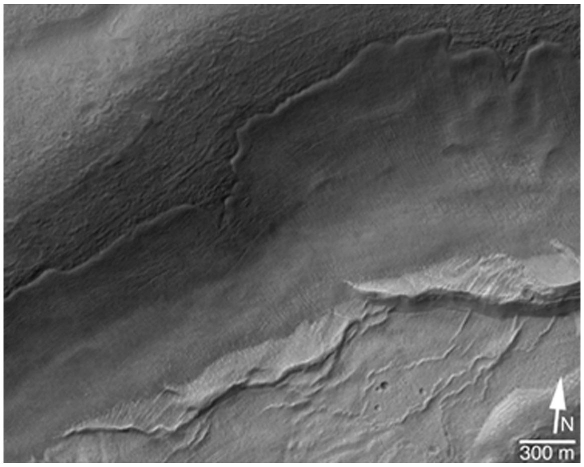


FIG. 13. A southeast-facing slope with a mantle of "pasted-on" terrain. This image is located near 38.1°S, 95.0°E in a depression at the head of Harmakhis Vallis. Material such as this has been interpreted by some (*e.g.*, Christensen, 2003) as a deposit of snow or ice beneath a residue of dust that is protecting the material from further sublimation. This is a subframe of MOC image S14-01956; sunlight illuminates the scene from the upper left.

the two types of landform are related globally, though Milliken *et al*. (2003) showed that there is a strong correlation between viscous flow features, dissected mantle terrain, and gullies within the ±30–50° latitude zones. Also, where gullies and mantles occur together, gullies cut the mantles and, in some cases, head at locations higher up the slope than the margin of the mantle/accumulation. It is possible that some gullies may be the final erosional product left by the melting of a prior mantle that is no longer present. In any case, our present understanding of whether there is a genetic relationship between gullies and mantles is missing some important details. Much work on the topic remains to be done.

Is there a part of Mars that currently has an active layer? Permafrost is ground that has remained frozen (temperature below the freezing point of water) for more than two consecutive years. An active layer in permafrost regions is defined as a near-surface layer that undergoes freeze-thaw cycles due to day-average surface and soil temperatures oscillating about the freezing point of water. A "dry" active layer may occur in parched soils without free water or ice, but significant geomorphic change through cryoturbation is not produced in these environments.

We have enough information to be able to conclude that a wet active layer is currently absent on Mars. Kreslavsky *et al*. (2006) used recent calculations on the astronomical forcing of climate change to assess the conditions under which an extensive active layer could form on Mars during past climate history. Their examination of insolation patterns and surface topography led them to conclude that an active layer should have formed on Mars in the geological past at high latitudes as well as on pole-facing slopes at middle latitudes during repetitive periods of high obliquity in the geological past. They examined global high-resolution 1996 MGS Mars Orbital Laser Altimeter (MOLA) (see http://ltpwww.gsfc.nasa.gov/tharsis/mola.html) topography and geological features on Mars and found that a distinctive latitudinal zonality of the occurrence of steep slopes and an asymmetry of steep slopes at middle latitudes can be attributed to the effect of active layer processes. They concluded that the formation of an active layer during periods of enhanced obliquity throughout the most recent period of the history of Mars (the Amazonian) has led to significant degradation of impact craters, which has rapidly decreased the steep slopes characterizing pristine landforms.

However, their analysis indicates that an active layer has not been present on Mars in the last ~5 Ma (millions of years) and that conditions favoring the formation of an active layer were reached in only about 20% of the obliquity excursions between 5 and 10 Ma ago. Conditions favoring an active layer are predicted to be uncommon in the next 10 Ma. The much higher obliquity excursions thought to have occurred in the earlier Amazonian appear to have been responsible for the significant reduction in magnitude of crater interior slopes observed at higher latitudes on Mars.

FINDING. Because some of the "pasted-on"–type mantle has a spatial, and some suggest a genetic, relationship to gullies (which in turn are erosional features possibly related to water), the "pasted-on" mantle may be a special region. The mid-latitude mantle, however, is thought to be desiccated, with low potential for the possibility of transient liquid water in modern times. Because the "pasted-on" mantle and some kinds of gullies may have a genetic relationship, the "pasted-on" mantle is interpreted to have a significant potential for modern liquid water.

7D. Glacial Deposits

Low latitude. The topic of glaciation—even at equatorial latitudes—has been discussed and debated for more than 3 decades (*e.g.*, Williams, 1978; Lucchitta, 1981; Kargel and Strom, 1992a,b; Head and Marchant, 2003; Neukum *et al*., 2004; Forget *et al*., 2006). New spacecraft data and evidence from terrestrial analogs have provided insight into these types of deposits in the equatorial region of Mars. Head and Marchant (2003) have presented evidence that the large lobate deposits on the northwest flanks of the Tharsis Montes volcanoes might have resulted from the accumulation of ice and snow, its flow outward to form glaciers, followed by the cessation of ice accumulation, collapse of the glaciers, and the production of distinctive glacial deposits that remain today (Fig. 14). Similar deposits occur around the base of the Olympus Mons scarp (Fig. 14) and are interpreted by some to represent debris-covered glacial deposits (*e.g.*, Milkovich *et al*., 2006), by others as landslides (*e.g.*, Carr *et al*., 1977). Climate modeling shows that, during periods of higher obliquity, water is mobilized from the polar regions, carried by the atmosphere to the tropics, rises along the western flanks of Tharsis, and is preferentially deposited as snow and ice as the rising moist air is adiabatically cooled (Forget *et al*., 2006).

One of the major impediments to the glacial interpretation of these deposits in the past has been the lack of occurrence of eskers, drumlins, and other indications of classic wet-based glaciation (*e.g.*, Zimbelman and Edgett, 1992). Recently analyzed terrestrial analogs of the huge deposits at Arsia Mons (~180,000 km^2) (Head and Marchant, 2003) and Pavonis Mons (Shean *et al*., 2005) show that glaciers typical of polar latitudes on Earth (cold-based glaciers, where the glacier flows over permafrost and deforms internally to the glacier, rather than with melting at the base) are a more appropriate analog to these tropical features on Mars. Thus, even when these glaciers were forming on Mars several tens to hundreds of millions of years ago (Head *et al*., 2005; Shean *et al*., 2006), there was little to no melting associated with them. Although we cannot determine whether there may be some residual ice at depth in these deposits, it would certainly be below a thick sublimation till because of the deposit's age—residual shallow ice is highly unlikely.

Geomorphic features that may have involved liquid water in these regions include those interpreted by some investigators to be impact craters, pingos, mud volcanoes, and the possible basal

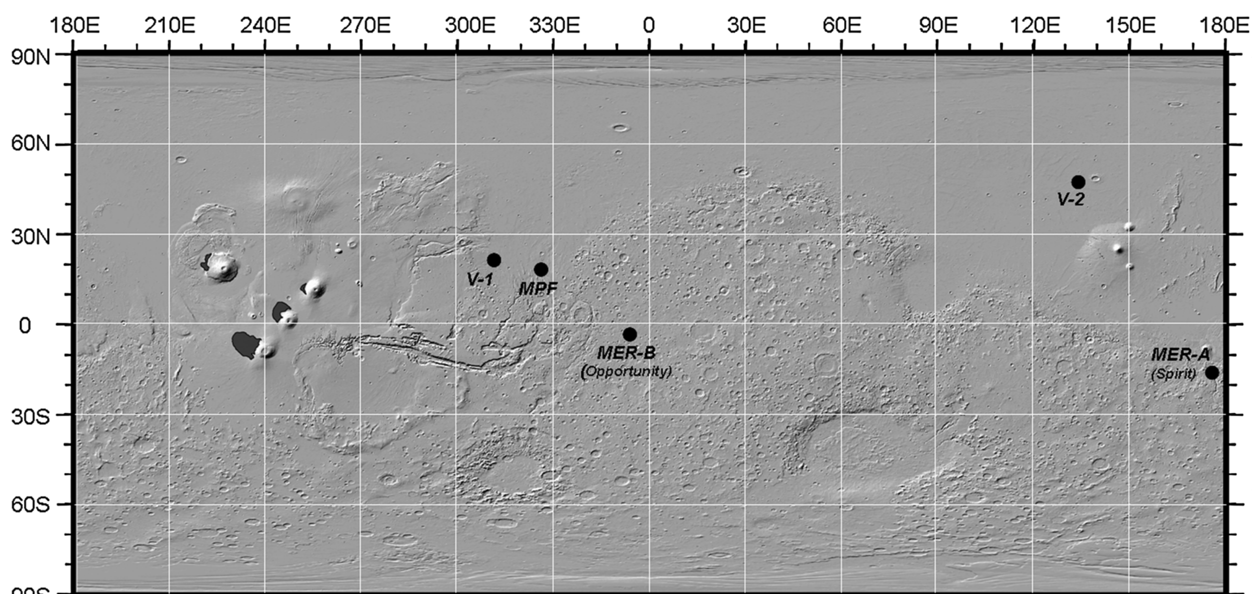


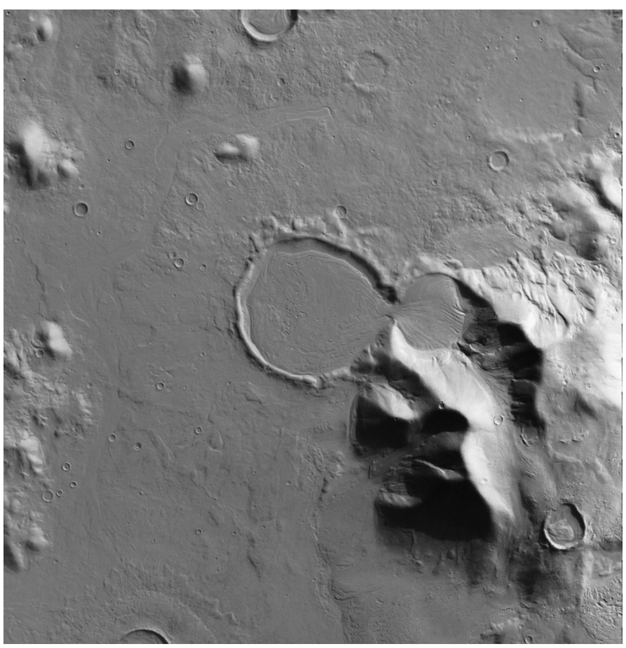
FIG. 14. Location of lobate deposits (black) on the northwest margins of the four big low-latitude volcanoes (Olympus Mons and the three Tharsis Montes). The map unit is from Scott *et al*. (1986–87) and was subsequently interpreted by Head and Marchant (2003), Shean *et al*. (2005), and Milkovich *et al*. (2006) to be tropical mountain glaciers. Lobate debris aprons and lineated valley fill are concentrated in the 30–50° north and south latitude bands (*e.g.*, Squyres, 1979; Squyres and Carr, 1986). (For base map details, see legend to Fig. 10.)

melts from ice caps. However, under current climatic conditions at middle to high latitudes where the supply of water from a shallow source may be present, the water source would be quickly exhausted, as surface recharge is virtually impossible.

Middle and high latitude. The mid-latitude regions exhibit additional geomorphic features for which some researchers have suggested that ice was or still is a part of the material. These include kilometer-scale flow features (*e.g.*, Fig. 15), interpreted by some to indicate they flowed in a viscous manner, which occur on some pole-facing slopes (Milliken *et al.*, 2003); aprons surrounding massifs and mesas in the Deuteronilus, Protonilus, and Promethei Terra regions; and lineated valley floor materials and concentric crater floor features that exhibit distinct erosional textures and morphologies that have been a topic of discussion since the Viking era (*e.g.*, Squyres, 1979; Squyres and Carr, 1986; Zimbelman *et al.*, 1989; Carr, 2001; Pierce and Crown, 2003; Berman *et al.*, 2005). The lineated valley floor material was once considered to be the product of creep or flow of ice-rich material (*e.g.*, Squyres, 1979). Although Malin and Edgett (2001) noted that lineated floor materials also occur in completely enclosed troughs, from which no material can flow, Head *et al*. (2006a,b) have shown that there is ample evidence for accumulation zones, zones of convergence and folding, and ablation zones very similar to terrestrial debris-covered valley glaciers and glacial land systems on Earth. Superposed craters, however, indicate that these deposits have not been active for tens to hundreds of millions of years (*e.g.*, Mangold, 2003). Although these features are typically located in the 30°–50° latitude range, effective maps are not yet available.

While there is no consensus about middle or equatorial latitude ice (*e.g.*, discussions of glaciers and other flows at middle latitudes), the presence of high-latitude ice on Mars, particularly in the north polar cap and associated with the south polar residual cap, is unquestioned. The key issue for the purpose of biological propagation is whether any of it ever exceeds a temperature of −20°C. In the high-latitude areas, we see no geo-

a



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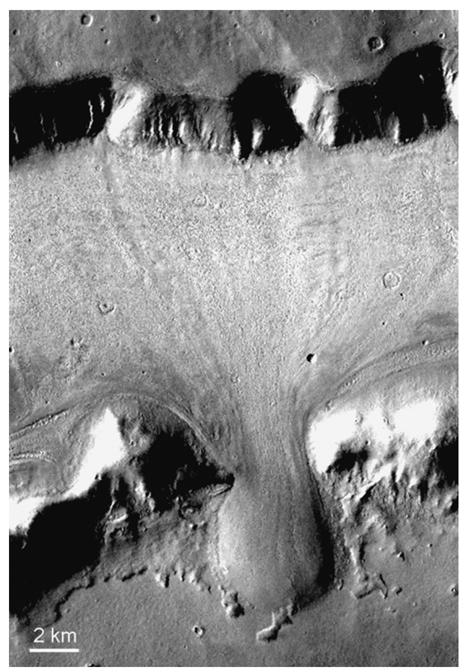

FIG. 15. (**a**) An example of a feature hypothesized to be a glacial deposit (Head *et al.*, 2005), located in Promethei Terra at the eastern rim of the Hellas basin, at about latitude 38°S and longitude 104°E. From HRSC (credits: ESA/DLR/FU Berlin (G. Neukum)]. Available at: http://www.esa.int/SPECIALS/Mars_Express/SEMN3IRMD6E_0.html. **(b)** A second example of a feature hypothesized to be a glacial deposit (image provided by Michael Carr). A lobate flow that appears to have been funneled between a gap between two obstacles in the fretted terrain of the Deuteronilus Mensae region (40°N, 25°E, THEMIS V12057009). The flow has been interpreted by some to be ice-rich, although this interpretation is not unique. From the superposed craters it is estimated to be tens to hundreds of millions of years old.

morphic evidence for melting, and there are theoretical grounds for believing that is not possible for melting to occur at present.

7E. Craters

Description, distribution, and age. Impact events can excavate to depths where ice and/or liquid water may exist, heat the surrounding region for a significant time, and create an environment that is out of thermodynamic equilibrium with its planetary setting. Although cratering events throughout martian history have certainly exceeded the threshold conditions for biological propagation, the heating is transitory, and the thermal anomalies are erased with time. Since craters are everywhere on Mars, the challenge is to identify which, if any, have the potential to retain enough heat to harbor, at present, liquid water within 5 m of the surface. Impact craters occur randomly in time and location on the martian surface.

Water in the target volume. The first question to address is which craters accessed volatile-rich material in the surface. Impact craters display a relationship between their depths of excavation (d) and their diameters at the time that maximum depth is reached (*i.e.*, the transient crater diameter, D_t). For small bowl-shaped craters ["simple craters"; typically $<$7 km diameter on Mars (Garvin *et al*., 2000)], the observed crater diameter (D) is approximately the transient crater diameter, and the excavation depth is approximately one-fifth of this diameter ($d \approx D/5$). For larger ("complex") craters, $d \approx D_t/10$, but the currently observed rim diameter is larger than the transient diameter primarily because of wall collapse (Melosh, 1989). Several empirically derived relationships between D and D_t for complex craters have been suggested. One example is that described by Croft (1985):

$$D_t = D_{sc}{}^{0.15}\, D^{0.85}$$

where D_{sc} is the simple-to-complex transition diameter of $\approx$7 km on Mars. Using this relationship, we find that

$$d \approx 0.13\, D^{0.85}$$

Fresh impact craters on Mars often display a layered (*i.e.*, "fluidized") ejecta morphology (Fig. 16), which is widely believed to form by vaporization of subsurface volatiles during crater formation (see review in Barlow, 2005), though laboratory experiments suggest that similar morphologies could be produced simply through the interaction of an expanding ejecta curtain and the atmosphere (*e.g.*, Schultz and Gault, 1979). The current best models for layered ejecta blanket formation suggest that the surface material must contain at least 15–20% ice before the layered ejecta patterns begin to form (Woronow, 1981; Stewart *et al*., 2001). Crater studies provide information about the entire history of a region, whereas GRS provides a snapshot of where the ice is today (within the upper meter of the sub-

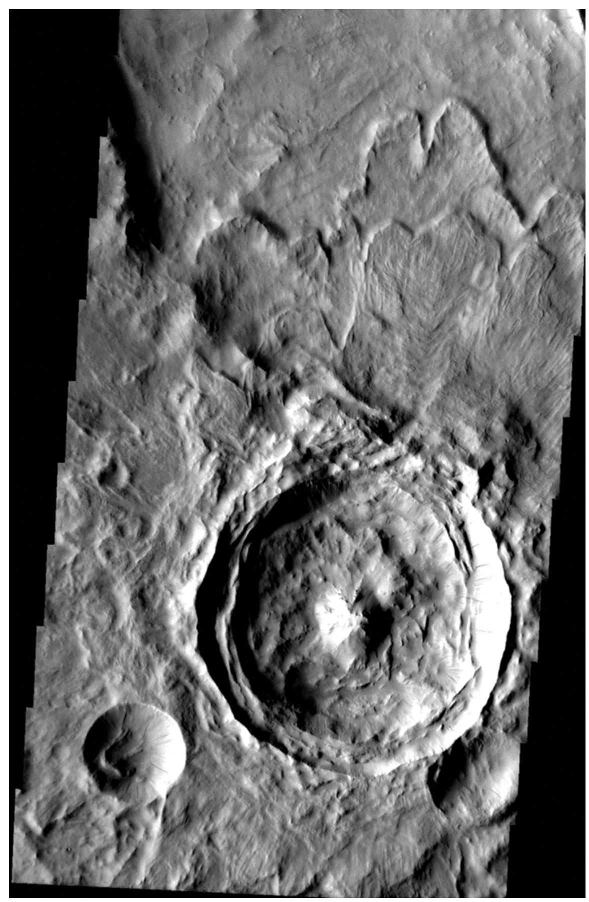

FIG. 16. One of the freshest large craters. This crater is 11.5 km in diameter, is located at 13.70°N 29.52°E, and shows a well-developed central peak (where models indicate that a long-lived hydrothermal system could operate). The adjacent craters are all covered with ejecta deposits, indicating those craters are older. The roughness of the floor is due to wall collapse during formation. Note the well-developed fluidized ejecta blanket to the right. V10297019 (THEMIS daytime visible); north is to the right.

surface). Thus, we do not expect these two approaches to agree in detail.

- Although regional variations occur, the smallest craters that show a layered ejecta morphology in the $\pm 30°$ latitude zone are typically between 3 and 5 km in diameter, which corresponds to excavation depths of 600 m to 1 km. This suggests that the uppermost martian crust in this latitude zone has been largely devoid of ice over the course of martian geologic history. Is there a systematic change in the onset depth as a function of time? Barlow (2004) argued, using ejecta characteristics, that there is no long-term variability in volatile concentrations at the depths encountered by these craters. However, recent HRSC data (Reiss *et al.*, 2005) suggest that there may be an indication of an increase to the depth of the volatile-rich layer over time. These crater-related observations are independent of assumptions about interaction between the atmosphere and the deep volatile reservoirs.
- In the ~40–60° latitude range, the following observations suggest that ice-rich material has been close to the surface in the past, a conclusion broadly consistent with modern GRS data: Onset diameters for single layer craters are <1 km in diameter; pedestal craters display even smaller diameters, and double layer craters are also suspected at these small crater diameters (resolution makes it difficult to distinguish single layer from double layer at these small diameters). However, research on craters between 50 and 500 m in diameter in the lineated terrain at 45°N suggests that ice is not currently present near the surface (McConnell *et al.*, 2006). Note that older maps of crater onset diameters using low-resolution Viking data (200 m/pixel resolution) give somewhat different results than maps based on higher-resolution MOC, THEMIS, and HRSC imagery.

Crater-induced heating and hydrothermal systems. The amount of heating associated with the impact-induced shock wave depends on the transient crater diameter and the location within or outside of the crater (Pierazzo *et al.*, 2005). Part of the kinetic energy associated with incoming bolides is converted into impact melt, and part of it goes into heating the target material during crater formation (Ivanov and Deutsch, 1999). Numerical modeling indicates that this heating is highest under the transient crater floor, within the region where central peaks and central pits form. In martian craters larger than 30 km, another significant heat source is the uplifted geotherm under the central uplift (Abramov and Kring, 2005a,b). Hydrothermal systems are expected to be particularly active within the central peak/central pit regions and in association with impact melts along the crater floor. Heat within the ejecta blanket and near the crater rim dissipates more rapidly, particularly with the cold surface temperatures prevailing on present-day Mars.

Water supply is a critical issue for the formation of an impact-generated hydrothermal system. Even small impact craters emplaced in ice can produce liquid water (Stewart *et al.*, 2001). Using the above-mentioned relationship between crater diameter and depth of excavation for areas with ice at a depth greater than 600 m, craters up to 3 km in diameter, where basement uplift is not an issue, are unlikely to produce a hydrothermal system of any extent or duration. Larger craters could potentially access liquid aquifers at depth.

Numerical models also indicate that impact melt sheets and central uplifts/pits associated with larger complex craters can produce active hydrothermal systems that can survive on time scales of up to ~10^6 years (Ivanov and Deutsch, 1999; Newsom *et al.*, 2001; Rathbun and Squyres, 2002; Abramov and Kring, 2005a,b). The lifetime of an impact-induced hydrothermal system is related to crater size and the geothermal gradient. The maximum potential lifetime for a hydrothermal system is the conductive cooling time for the hot rock produced by an impact. For example, a recent calculation by Abramov and Kring (2005a) has estimated a 67,000-year lifetime for a hydrothermal system in a 30-km-diameter crater on Mars.

Based on the above general considerations, SR-SAG proposes the conservative guidelines in Table 4 for the maximum amount of time a crater

TABLE 4. APPROXIMATE RELATIONSHIP BETWEEN CRATER SIZE AND MAXIMUM DURATION OF HEATING

Crater size (diameter)	*Time for which crater environment has potential to retain enough heat to exceed threshold conditions*
3 km	100 years
10 km	1,000 years
30 km	100,000 years

environment has the potential to retain enough heat in the shallow martian subsurface to exceed the threshold conditions for microbial propagation, assuming a supply of water is available. In general, the thermal lifetimes for craters smaller than 3 km in diameter are too short to be significant for the purpose of special region analysis. In addition, very young and fresh craters larger than 30 km in diameter have not been identified (Table 5).

Cratering frequency. A 10-km-diameter crater is expected to form on Earth approximately every 10^5 years (Morrison *et al*., 1994). The impact rate on Mars is considered to be ~1.3 times higher than the terrestrial impact rate because of Mars' proximity to the asteroid belt, but the impact energy will be less because the typical asteroid impact velocity is 8.6 km/s on Mars compared with ~17 km/s on Earth. Adjusting the terrestrial impact probability plot for Mars suggests that one 40-km-diameter crater could be expected to form approximately every million years on Mars, and a 10-km-diameter crater would statistically form every ~30,000 years. Thus, although the probability is very low that a crater of sufficient size has formed recently enough to retain an active hydrothermal system to the present, there is a small chance that such a crater exists.

Identification of the most recent large craters. Young craters can be identified by the following characteristics:

1. Sharp rim and crater depth approximately equal to those values expected for a pristine crater of equivalent size.
2. No superposed features on either the crater or its ejecta blanket. Such superposed features would include dunes, floor deposits, tectonic/fluvial features, or small impact craters.
3. Ejecta blanket and interior morphologies that are sharp and well preserved.
4. The crater and its ejecta blanket displaying thermally distinct signatures in daytime and/or nighttime infrared views.

Very few impact craters >5 km in diameter display these characteristics. Table 5 lists northern hemisphere craters identified from MOC, MOLA, and THEMIS analysis, which appear to be among the freshest craters on the planet (from data of Malin and Edgett, 2001; Smith *et al*., 2001; Christensen *et al*., 2004; interpretation by Barlow, 2004). These craters have a similar rating of extremely fresh using the criteria listed above. It should be note that none of these craters is named, and this list will clearly keep growing as our exploration of Mars proceeds. We do not have precise ways of dating craters on Mars; however, the approximate age can be estimated from the degree of degradation of the crater morphology (Barlow, 2004). Based on this kind of analysis, we do not have reason to believe that any of the craters in Table 5 are as young as the limits specified in Table 4.

FINDING. No craters with the combination of size and youthfulness to retain enough heat to exceed the temperature threshold for propagation have been identified on Mars to date.

7F. Young Volcanics

Description and distribution. Volcanism or magmatic intrusion may be capable of generating spe-

TABLE 5. FRESHEST LARGE CRATERS IN THE NORTHERN HEMISPHERE OF MARS IDENTIFIED TO DATE

Latitude (N)	*Longitude (E)*	*Diameter (km)*	*Central structure*	*Ejecta*[a]
7.03	117.19	18.0	Central peak	MLERSRd
7.16	174.41	9.6	Floor pit	MLERS
8.93	43.82	10.9	Summit pit	MLERS
12.10	169.24	5.9	Summit pit	SLERS
13.70	29.52	11.5	Central peak	MLERS
16.95	141.70	13.6	Floor pit	MLERSRd
19.51	141.18	9.2	Floor pit	MLERS
20.01	246.68	7.9	None	SLERSRd
23.19	207.76	28.3	Central peak	MLERSRd

[a]Ejecta classifications from Barlow *et al*. (2000).

cial regions on Mars by warming near-surface rocks and melting ground ice. Some key examples of fluvial channels in spatial and temporal association with volcanic eruption products have been documented on Mars (*e.g.*, Mouginis-Mark, 1990; Burr *et al.*, 2002; Mouginis-Mark and Christensen, 2005). Subsurface magmatic intrusions (*e.g.*, dikes and sills) have also been implicated in fluvial activity in volcanic regions (*e.g.*, Tanaka *et al.*, 1998; Head *et al.*, 2003b). Since the time period of interest for special regions consideration is from now until 100 years (from the date of a mission's arrival) in the future, we need to consider recent volcanic rocks that might still be warm and the possibility of future volcanic eruptions.

A first-order assessment can be made from global geologic mapping of Mars and model-dependent absolute ages derived from impact crater densities. The youngest volcanic materials on Mars are considered to have formed during the Late Amazonian Epoch, in which the number of craters per 10^6 km^2, or $N(1)$, is less than 160 (Tanaka, 1986). These Late Amazonian volcanic materials were originally mapped with Viking images at 1:15,000,000 scale (Scott *et al.*, 1986–87). That mapping effort led to identification of relatively young volcanic rocks associated with the Tharsis Montes and Olympus Mons. In addition, young deposits of the Medusae Fossae Formation have an uncertain origin but could be volcanic (*e.g.*, Scott and Tanaka, 1982). Improved mapping at 1:15,000,000 scale, based on MOLA topographic data and early THEMIS data, covers the northern plains of Mars, or about a third of the planet. Based on this mapping, the youngest volcanic province on Mars is map unit AEC$_3$ from Tanaka *et al.* (2005), the distribution of which is shown on Fig. 17. These materials include a broad, lightly cratered region of lava flows and vents associated with Cerberus Fossae (Plescia, 1990, 2003; Keszthelyi *et al.*, 2000), though some investigators have proposed that the materials are not volcanic (Murray *et al.*, 2005) or, if volcanic, may be exhumed and therefore not necessarily young (Malin and Edgett, 2001).

Magmatic intrusions would provide additional sources of near-surface heat. Basaltic eruptions are generally emplaced by dikes about 1 m wide on Earth, but may exceed 100 m in width in association with flood volcanism (*e.g.*, Wada, 1994). Intermediate and silicic magma bodies on Earth may amount to several times the volume of extrusive products (*e.g.*, Tanaka *et al.*, 1986). Thus vent areas of silicic and basaltic flood lavas on Mars may remain warm for particularly long periods of time and perhaps generate geysers and

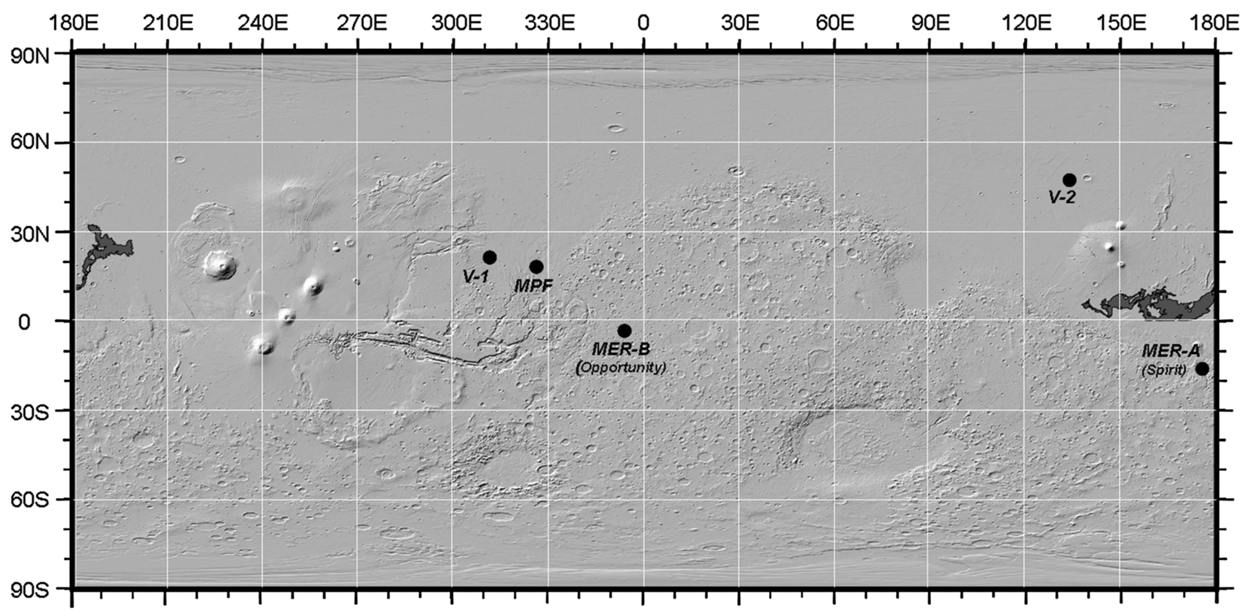

FIG. 17. Shaded-relief map of Mars showing the distribution of the youngest volcanic rocks on Mars [map unit AEC$_3$ from Tanaka *et al.* (2005)]. These volcanics are of Late Amazonian age, which includes geologic time from the present to about 300–500 Ma. Although there is no current indication that any of these rocks are as young as 1,000 years, we cannot say with certainty that the activity in this volcanic region has ceased, and that future eruptions are impossible. (For base map details, see legend to Fig. 10.)

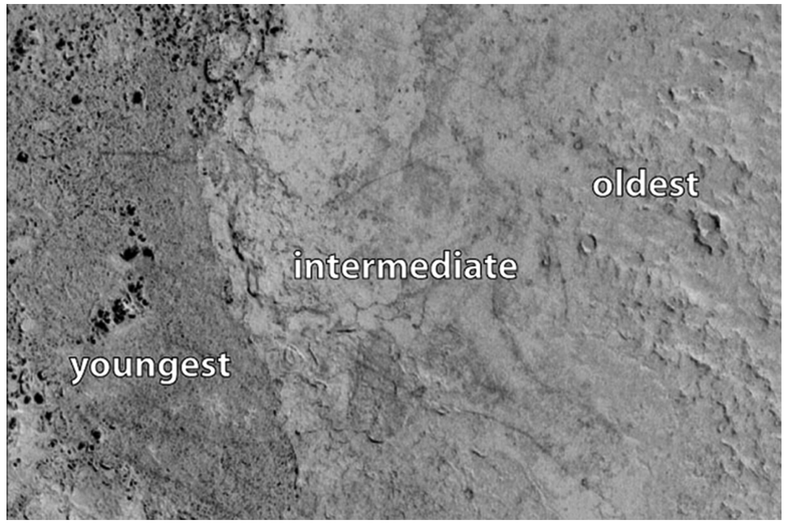

FIG. 18. Three distinct, apparent volcanic surfaces in Elysium Planitia progressively younger from right to left (appearing rugged and cratered, bright and partly incised, and dark, respectively). The youngest unit is Late Amazonian [in the map of Tanaka *et al*. (2005)], the middle unit is Late and Middle Amazonian, and the cratered unit is Late Hesperian (Cerberus Fossae 3 and 2 and Utopia Planitia 2 units, respectively). This shows the episodic nature of eruptions in volcanic regions. Image width 3.7 km; part of MOC image R11-01377.

hydrothermal springs where water recharge occurs. Such activity can prevail for as long as hundreds of thousands of years for particularly large, shallow magma bodies. For these reasons, any volcanic vent regions that display evidence for anomalously warm surface temperatures should be considered special regions. However, investigations of surface temperatures from MGS Thermal Emission Spectrometer and Odyssey THEMIS data have not revealed any such locales on Mars thus far.

Age. Volcanic products on Mars are largely considered to be mafic (*e.g.*, Greeley and Spudis, 1981, and many Mars meteorite studies), which are dark in color. The relatively youngest volcanic materials are, therefore, assumed to be relatively dark (see Fig. 18) because of the progressive accumulation of thin coatings of dust that collect after dust storms. Age can also be estimated from cratering density. Viking, MOC, HRSC, and THEMIS data demonstrate that no such dark, uncratered surfaces occur on Mars (*i.e.*, there are no black lava flows, like the ~900-year-old flows from Sunset Crater in northern Arizona). The highest-resolution images, from MOC, show the volcanoes and flows of the Tharsis Montes and Olympus Mons, for example, to have been covered by mantles of fine-grained material (*i.e.*, dust) that, in most cases, has been there long enough for the materials to become somewhat indurated and sculpted by wind erosion (Malin and Edgett, 2001). The youngest volcanic materials might be those associated with the Cerberus Fossae/southeastern Elysium Planitia and Marte and Athabasca Valles systems, south of the Elysium rise, but these, too, are mantled with dust.

The absolute ages of these possibly volcanic surfaces are estimated by crater density and models of crater production rate (Hartmann and Berman, 2000; Hartmann and Neukum, 2001). Such age determinations have to be considered with caution. First, there is debate at present as to whether secondary craters dominate the populations of subkilometer craters (McEwen *et al*., 2005), which have been incorporated when evaluating the ages of the youngest surfaces (Hartmann and Neukum, 2001). Second, the error in the model ages is uncertain and considered to be a factor of 2 (Hartmann and Neukum, 2001). Third, the geology of the dated rocks may be complex in some cases. If the crater counts include multiple outcrops of lava flows and other volcanic materials, then the counts provide a mean age. Some surfaces may be much older than others within the counted area. Crater obliteration processes, particularly for softer materials such as volcanic ash, may result in surface ages that are much younger than the rock emplacement ages.

In addition, no known "warm" surfaces—that is, those that can be attributed to volcanic or magmatic heating—have been detected on Mars from spacecraft imaging or thermal remote-sensing datasets.

The possibility of eruptions within the next 100 years. To establish a volcanic recurrence rate, information is needed for multiple eruptions. Geologic mapping and crater dating investigations, however, have not yet resolved this to any useful precision. For example, for the materials interpreted to be geologic units that consist of multiple, overlapping lava flows, it is uncertain in detail how many eruptive episodes were involved. Thus the duration of the Late Amazonian constrains the mean recurrence interval for volcanic materials of this epoch. Its duration ranges from 150 to 1100 Ma, given the factor of 2 age uncertainties (see Hartmann and Neukum, 2001) (Fig. 18). The recurrence interval is the inverse of the ages in years times 100 for a 100-year period, or 9.1×10^{-8} to 6.7×10^{-7}. If, however, the materials of the Cerberus Fossae region and other volcanic materials were to have a recurrence interval of 10 Ma, equivalent to a younger estimate of its age from subkilometer crater densities (Hartmann and Neukum, 2001), then the probability of recurrent volcanism for a 100-year period at any of these sites is on average $<10^{-5}$. Volcanic episodicities related to longer time periods have also been discussed by Wilson *et al*. (2001) and Neukum *et al*. (2004).

In summary, the recurrence and extent of volcanic eruptions on Mars appear to be sufficiently low in the most recently active regions that the occurrence of special regions due to either recent or renewed volcanism seems unlikely. However, without more precise geologic mapping of eruption deposits and their age dating, it is not possible to determine exactly the recent volcanic recurrence interval of local volcanic areas on Mars that may be of interest to future landed missions. For example, Athabasca Valles, which has been proposed by some to be one of the youngest fluvial and volcanic systems on Mars, was seriously considered as a candidate landing site for the MERs (*e.g.*, Golombek *et al*., 2003), yet its detailed geologic history remains elusive.

Potential to exceed propagation thresholds. Since volcanic heat is lost with time, only extremely young volcanics have the potential to exceed the propagation thresholds. An obvious question, therefore, is how much time is too much? A simple calculation was done using an infinitely thick slab with an initial temperature of 1000°C cooling in an environment with a surface temperature of −40°C, and the following parameters: thermal conductivity of the surface = 2.5 W/m/K, heat capacity of the surface = 800 J/kg/K, and density of the surface = 2,600 kg/m^3. As is typical of solutions to the heat flow equation, the initial cooling is very rapid, but after the first few tens of years, the cooling slows considerably. The temperature of the surface drops to less than −20°C within about 1,000 years. Although it is beyond the scope of this report to analyze the possible cooling histories of all configurations of eruptive volcanic rocks, the 1,000-year figure is a useful practical guideline. Volcanic rocks older than that cannot have retained enough heat to exceed the −20°C propagation limit within the upper 5 m of the martian crust. This limit is very conservative in the sense that the upper 5 m will almost certainly be dry as well as cold. It is more difficult to put a limit on the amount of time for loss of water from the system, but this is not necessary since both threshold conditions must be satisfied for propagation. More detailed treatments of the interactions of magma and ice on the Earth and Mars can be found in Wilson and Head (2002) and Head and Wilson (2002), respectively.

Volcanic rocks younger than 1,000 years old have not been discovered on Mars, so we do not have evidence for shallow volcanic-related special regions.

<u>FINDING</u>. **We do not have evidence for volcanic rocks on Mars of an age young enough to retain enough heat to qualify as a modern special region or suggest a place of modern volcanic or hydrothermal activity.**

7G. Slope Streaks

Description and age. Dark- and light-toned slope streaks (Fig. 19) occur in dust-mantled regions, particularly Arabia Terra, Tharsis, and various hills and slopes in the regions between Tharsis and Elysium in the northern hemisphere (Fig. 20). Sullivan *et al*. (2001), Schorghofer *et al*. (2002), and Aharonson *et al*. (2003) have presented basic summaries on the subject. Slope streaks typically originate at a point on a slope that is often associated with a roughness element, such as a small knob. The streaks then continue downslope for hundreds of meters or more. They usually have a narrow, well-contained shape that spreads out into a single fan or multiple, sometimes anastomosing or braided, forms as it proceeds down the slope. Slope streaks have sharp boundaries, usually (but not always) with a nearly constant brightness throughout the feature, and their albedo contrast with the surrounding terrain has been seen to decrease over time.

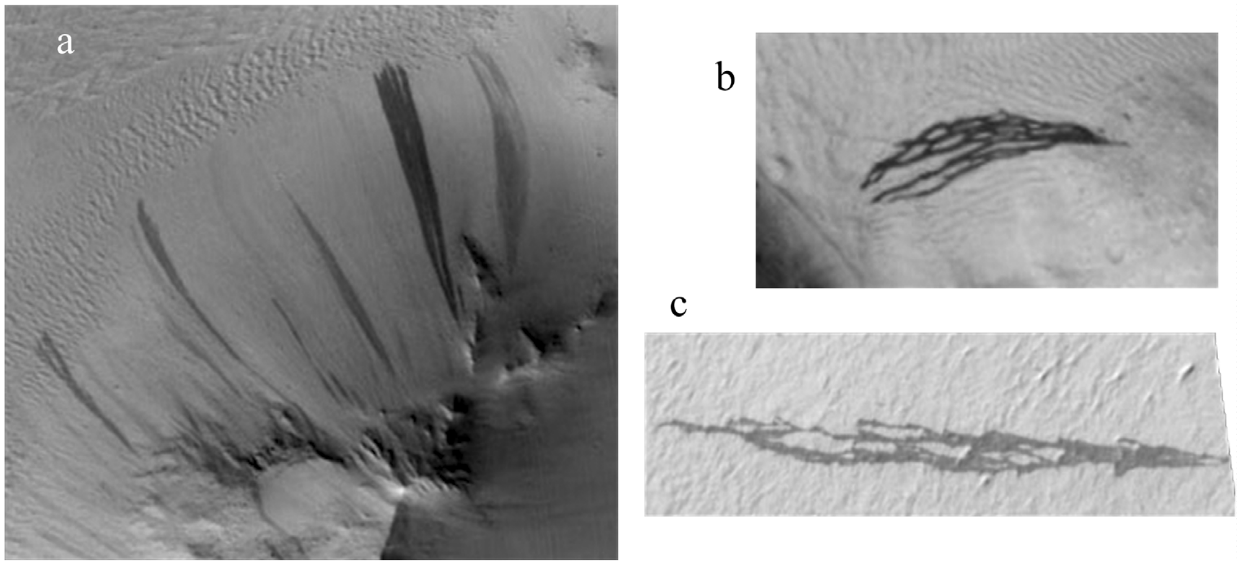

FIG. 19. Slope streak examples (from Phillips and Chyba, 2006). Typical slope streaks (**a**, MOC image M16-00596, 27°N 227°E, 5.9 m/pixel) have a point source and a wedge-shaped, sometimes digitate, appearance. More complex examples (**b**, image M16-00596, 27°N 227°E, 5.9 m/pixel; **c**, image E02-00308, 0°N, 195°E, 4.4 m/pixel) have an intricate, braided, sometimes anastomosing shape. These features are widely interpreted to be the result of dry dust avalanches, but one working hypothesis for what triggers them to form involves water.

Slope streaks are relevant to the present analysis regarding special regions because they form in response to modern geologic processes, and some investigators have proposed that these features might involve the action of liquid water (Ferris *et al.*, 2002; Schorghofer *et al.*, 2002; Miyamoto *et al.*, 2004b). New slope streaks have been identified in comparisons between Viking and MOC images, and new streaks have even been observed to have formed during the MGS mission, on time scales as short as ~100 days (Malin and Edgett, 2001; Sullivan *et al.*, 2001; Schorghofer *et al.*, 2002; Aharonson *et al.*, 2003; Miyamoto *et al.*, 2004b).

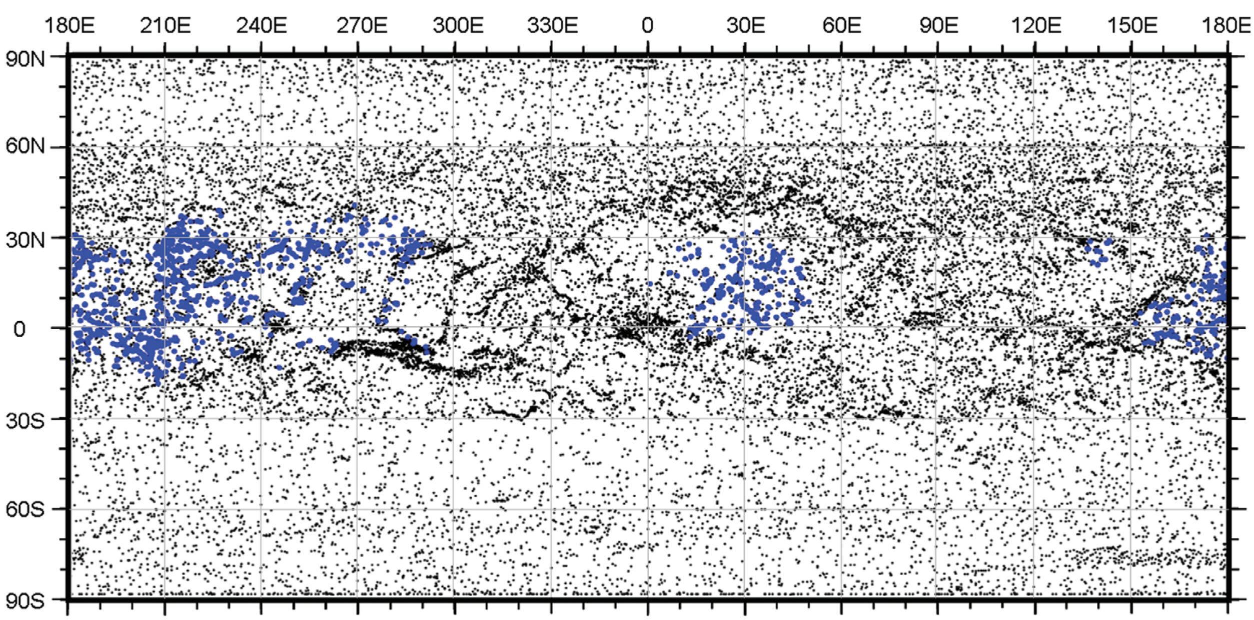

FIG. 20. Map of the distribution of dark slope streaks (from Aharonson *et al.*, 2003).

Possible relationship to water. There are two current working hypotheses for their formation mechanism:

- The dry models suggest that the streaks form through dust movement. In the dry process, oversteepening slopes caused by air fall deposits of dry dust eventually collapse, forming a dust avalanche (Williams, 1991; Sullivan *et al*., 2001; Miyamoto *et al*., 2004b). In support of the dry model is that all light and dark slope streaks occur in low thermal inertia areas mantled by accumulations of dust thick enough to be evident as such in MGS MOC images.
- The second class of models involves water, but the role of water varies between models. Early models such as that of Ferguson and Lucchitta (1984) suggested that the dark streaks could be stains on the surface produced by wet, briny debris flows. These flows could form when a slope intersected an aquifer, allowing the periodic release of fluid from a wet subsurface layer. Schorghofer *et al*. (2002) suggested that water could lubricate avalanches, or the sublimation of near-surface ice could trigger mass movements. Ferris *et al*. (2002) proposed a model involving groundwater springs that infiltrate and saturate the surface, creating the dark streaks. Miyamoto *et al*. (2004b) concluded that the streaks are not debris flows, but could not eliminate the possibility that water is involved in the streak-forming process.

Despite the various hypotheses in the literature, two key observations indicate that wind is a controlling factor for at least some slope streaks: (1) examples have been seen where dust devil tracks turn into a slope streak once a crater rim has been crossed; and (2) a population of slope streaks is clearly emplaced preferentially on west-facing slopes (Baratoux *et al*., 2006)—this correlates with wind direction. The abundance of slope streaks decreases sharply at about 33°N, and they are absent at higher latitude. Baratoux *et al*. (2006) interpreted this as an indication that poleward of 33°, the ubiquitous dust is ice-cemented and, thus, not amenable to flow, whereas equatorward of 33° the dust is desiccated and flows easily. For these reasons, it is unlikely that substantial quantities of liquid water are involved in slope streak formation. There is, however, an open question on the role of trace amounts of transient H_2O in triggering these dust avalanches, for example, by changing cohesion or porosity of the dust matrix.

7H. Recent Outflow Channels?

Description and age. One of the prominent aspects of martian geology is the evidence for ancient outflow channels. Could the process that created them continue to the present era? The youngest such feature that has been dated is Athabasca Valles (10°N, 157°E)—Burr *et al*. (2002) have given the age of the latest flood at Athabasca as 2–8 million years. Some other channels nearby are only slightly older. Furthermore, there are also some young channels just to the southeast of Olympus Mons (Mouginis-Mark, 1990) that Basilevsky *et al*. (2006) have dated at 20 million years. These examples indicate the presence of deep groundwater (or ice) that might have episodically erupted to the surface along faults (or caused by dikes). Although this type of feature may have continued to form until the geologically recent past, we do not have evidence that any are active today or have reason to predict that any will be active within the next 100 years.

7I. The Nondiscovery of Geothermal Vents

An important objective of the THEMIS infrared investigation on Mars Odyssey has been the search for temperature anomalies produced by surface cooling or heating due to near-surface liquid water or ice, or hydrothermal or volcanic activity. THEMIS has mapped virtually all of Mars at night in the infrared at 100 m/pixel resolution and observed portions of the surface a second time up to 1 Mars year later. An analysis has been performed of all of these images to search for maximum temperatures greater than those expected from rocks or bedrock alone (220K), and no example has been found in any image of a temperature that requires an internal heat source (P.R. Christensen, personal communication, 2006).

The THEMIS database has also been searched for spatial patterns that might indicate evaporative cooling associated with near-surface water or recent volcanic activity or hydrothermal heating. The THEMIS nighttime images exhibit very high spatial variability owing to variations in surface properties such as rock layers, exposed bedrock, and sand and granule dunes (P.R. Christensen, personal communication, 2006). These thermal variations greatly complicate the search for anomalous patterns associated with subsurface water. To overcome the effects of these surface temperature complexities, a search for seasonal changes in surface temperature was undertaken in an effort to isolate changes due to dynamic heating or cooling processes from those due to thermophysical properties. This process involved the use of histogram adjustments of the overlapping portions of two images to remove the additional complicating temperature changes produced by the 2-h variation in local time of the Odyssey orbit, and by the differences in season between the different observations. In addition, numerous image pairs have been converted to thermal inertia using the THEMIS Standard Thermal Model as a quantitative means to remove the time of day and seasonal effects. To date, no significant changes have been detected, but this analysis will continue throughout the Odyssey extended missions. Based on the analysis thus far, there is no evidence in the THEMIS thermal images for the existence of near-surface liquid water or ice that is close enough to the surface to be capable of producing measurable thermal anomalies.

The MGS MOC investigation has also addressed this issue by focusing on imaging of relatively young volcanic and impact landforms, and searching for evidence that activity is occurring today—such as eruption, production of flows (mud or lava), and the like. No such features have been found in MOC or other orbiter images.

FINDING. **Despite a deliberate and systematic search spanning several years, no evidence has been found for the existence of thermal anomalies capable of producing near-surface liquid water.**

7J. The Possibility of Low-Latitude Ground Ice

The Mars Odyssey GRS data (Fig. 5) clearly reflect high-latitude ice in the vicinity of both poles, but also some significant positive equatorial anomalies in places like Arabia Terra. A number of authors have concluded, in large part based on the instability of ice within a meter of the surface (the volume within which the GRS instrument can detect ice), that these anomalies are related to hydrated minerals (*e.g.*, Basilevsky *et al.*, 2003). For example, the high hydrogen level in low latitudes is entirely consistent with the observations of hydrated minerals found by MER and by OMEGA (the visible and infrared mineralogical mapping spectrometer on Mars Express), and there is no evidence of layering of hydrogen in the low-latitude anomalies.

However, there has also been some discussion of the possibility that these anomalies represent low-latitude ground ice. For example, it has been suggested that, when the south polar cap has exposed water ice, the water vapor content of the martian atmosphere could climb and might make ground ice stable to low latitudes (*e.g.*, Jakosky *et al.*, 2005a,b). It may be possible for ice to be emplaced onto the surface or into the top meter during these episodes. While the ice is not stable today, it has been suggested that ice that had been previously deposited might still exist in a transient state while in the process of disappearing.

However, diffusion processes that control potential ice deposition in the equatorial regolith are thought to occur on 100–5,000 year time scales (Feldman *et al.*, 2005). If the time scale of uncovered polar water ice is shorter than 100 years, it is difficult for significant amounts of water ice to accumulate. Polar cap exposure on longer time scales, such as would result from orbital oscillations, might result in higher concentrations of equatorial ground ice, but also would result in a different distribution of high-latitude ground ice than is observed today (Mellon *et al.*, 2004). In addition, "transient" ice would be stable and in equilibrium with the atmosphere only in regions and depths where subsurface temperatures are typically below the new frost point (for even a 10-fold increase in atmospheric water the new frost point would only be 212K). Subsequent replenishment of the CO_2 cap on the south polar region and subsequent reduction of atmospheric water vapor would result in removal of now unstable ice by sublimation on the same time scale as its emplacement. Thus, presently observed equatorial concentrations of hydrogen are more likely to be the result of one or more hydrated minerals and not due to transient water ice (Feldman *et al.*, 2005).

The possibility of massive subsurface ice deposits. There has also been discussion in the literature of the possibility of massive subsurface ice deposits on Mars, even at low latitude, that represent the remains of a frozen ocean (postulated by Clifford, 1993; possibly observed by Murray *et al.*, 2005). The site described by Murray *et al*. (2005) is in southern Elysium (around 5°N latitude and 150°E longitude), where they reported geomorphic evidence consistent with a presently existing frozen body of water, which they interpreted as the remains of surface pack-ice. The ice slab is interpreted to have an age from crater counts of 5 ± 2 Ma, an original mean thickness of 45 m, and a present-day mean thickness of 30 m (with a significant part of the difference representing an overlying sublimation lag). The essential question, for the purpose of this analysis, is not whether such interpretations are correct, but whether, if they are correct, this could constitute a niche environment in which the biological thresholds for T and a_w could be exceeded within the upper 5 m. As discussed in Sec. 6 of this report, the ice itself is not intrinsically special, and if in fact it is preserved at all in the shallow subsurface, it would be in a place where the temperature is low enough that a_w is below the biological threshold.

7K. The Polar Caps

In the initial COSPAR definition of special region (see DEFINITION #1), the polar caps are mentioned as an example of a special region. Using findings of this analysis and the proposed guidelines in this report, the polar caps would no longer fit the definition nor serve as valid examples. Jakosky *et al*. (1993) (as well as others) have calculated that, at the present obliquity, the maximum summer temperatures of the ice surface typically reach about 200K at the pole (the north and south poles are similar in this respect) (Fig. 21). This temperature is consistent with the amount of atmospheric water vapor observed in the atmosphere during northern polar summer, which is about 10 precipitable microns with a 10 km scale height. More water vapor is observed in the atmosphere during the north polar summer (~100 precipitable microns), which is consistent with a surface ice temperature of ~210K.

Contributing to the perpetual low temperature is not only the latitude (hence low sun angle) but also the high conductivity of solid ice. Polar ice is highly conducting material (*e.g.*, Paige and Keegan, 1994). Paige and Keegan (1994) determined the thermal inertia of the polar ice caps and found them to be in the neighborhood of 2,000 (meters-kilograms-seconds units), which is consistent with solid ice. While snow and firn might produce more insulating material, models and measurements of thermal inertia (Arthern *et al*., 2000) indicate that the polar ice is, indeed, highly conducting, and the annual skin depth is around 10 m. That means that seasonal changes must heat up a slab 10 m thick to affect a significant change in surface temperature. The argument is particularly compelling at the north pole, which currently experiences cool summers at aphelion (and will continue to do so for many thousands of years). The south polar cap, despite receiving more summer sunlight, is protected by a layer of highly reflective CO_2 ice, which holds the surface temperature at a constant 145K. While these models include the high conductivity and heat capacity of ice, none of these models takes into account the additional latent heat loss when the sublimation rate is high—this would cause the present models to be overly optimistic regarding the potential for warming the polar cap material.

The warmest places on the polar caps might be steep slopes where the strata of the polar layered deposits are visible and the ice may be actively retreating. Equatorward-facing slopes or

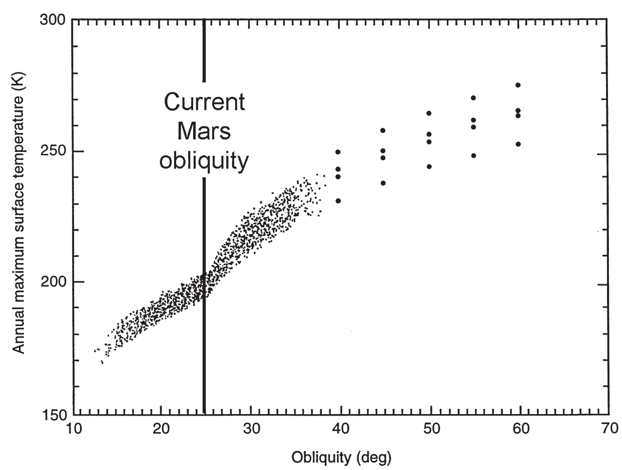

FIG. 21. Maximum summer time temperature of the ice surface at the pole (similar north and south). The scatter is due to eccentricity and perihelion variations. The largest surface temperature is roughly 205–210K. From Jakosky *et al*. (2003), using models in Jakosky *et al*. (1993).

low-latitude ice would indeed receive more insolation. This may account for the higher atmospheric water vapor if the slopes reach temperatures of 210–212K.

Even local heating of the surface of this ice to temperatures greater than −20°C is not thermodynamically possible in today's climate. As described above, Mars' large mass of polar ice is at an annual mean temperature of approximately −110°C (at least 50°C lower than the coldest ice on Earth) and is known to have high thermal conductivity. Consider, for example, local warming due to a patch of dark particulate material [analogous to the source of cryoconite holes on Earth (NRC, 2006, pp.75–78)]. Assuming albedo of 0.15, nominal thermal properties, and 85°N latitude, model calculations indicate that the ice temperature will never exceed −43°C. Alternatively, consider the heat balance required to maintain ice at −20°C. At this temperature, radiation to the sky would be ~220 W/m^2, and the latent heat of sublimation would be nearly 60 W/m^2, for a total of 280 W/m^2. Conductive losses to the subsurface would be of comparable magnitude, depending on the specific model. But even with a perfectly clear sky and a low (dark) albedo of 0.15, a flat surface at 85°N in the peak of the day in midsummer only receives ~210 W/m^2 insolation.

FINDING. The martian polar caps are too cold to be naturally occurring special regions in the present orientation of the planet.

8. REVISION OF THE SPECIAL REGION DEFINITION AND GUIDELINES

SR-SAG concluded that it could fulfill its assignment concerning reducing ambiguity in the term "special region" by retaining the original COSPAR definition and adding an updated set of clarifications and implementation guidelines (DEFINITION #2). (Spacecraft-induced special regions are mentioned in DEFINITION #2 but are developed in Sec. 10.) These guidelines are as quantitative as possible based on the SR-SAG analysis, and are intended to allow the definition to be interpreted in a common way by the many different interest groups who are stakeholders in Mars exploration.

DEFINITION #2:

Existing definition of "special region" with proposed implementation guidelines:

A special region is defined as a region within which terrestrial organisms are likely to propagate, *or* a region that is interpreted to have a high potential for the existence of extant martian life forms.

Proposed implementation guidelines:

1. *Definitions*. For the purpose of this definition, *propagate* means to reproduce. Other kinds of activity, including cell maintenance, thickening of cell walls (as aspect of growth), and mechanical dispersal by aeolian processes are not sufficient.
2. *Period of applicability*. The time period over which these guidelines are to be applied is defined as from the present until 100 years after spacecraft arrival on Mars.
3. *"Non-special" regions*. A martian region may be categorized non-special if the temperature will remain below −20°C ***or*** the water activity will remain below 0.5 for a period of 100 years after spacecraft arrival. All other regions on Mars are designated as either special or uncertain.
    a. *Uncertain regions*. If a martian environment can simultaneously exceed the threshold conditions of −20°C ***and*** a_w over 0.5, propagation may be possible. It may not be possible to show that such environments are capable of supporting microbial growth, but such areas will be treated in the same manner as "special regions" until they are shown to be otherwise.
4. *Induced special regions*. Even in an otherwise "non-special" region, a spacecraft may create an environment that meets the definition of a "special" or "uncertain" region, as described above. Because of the many dependencies related to spacecraft design, planned or accidental operations, or landing site, the possibility of a mission causing a spacecraft-induced special region should be analyzed on a case-by-case basis.
5. *Impact scenarios*. As a practical consideration for evaluating accidental impact scenarios involving both naturally occurring and induced special regions, it is considered sufficient to consider maximum crater depth to be <5 m for impacting hardware of <2,400 kg.

9. DISCUSSION OF NATURALLY OCCURRING SPECIAL REGIONS

9A. Risk Acceptability

A key term in the definition of special region (see DEFINITIONS #1 and #2) is "*likely* to propagate." "Likely" implies a probability. The lay usage of the word "likely" is that it implies a probability level of 50%. However, that is clearly not the intent of COSPAR for this application—something significantly lower, but non-zero, would be consistent with probability thresholds used elsewhere in PP policy. According to this definition, not every martian environment that has non-zero probability to exceed the threshold conditions for propagation is special; only those for which the magnitude of that potential reaches "likely" would qualify.

Risk consists of a probability and an adverse consequence (for example, the propagation of terrestrial organisms placed in a certain martian environment). A crucial question is the degree to which this risk is acceptable. SR-SAG considered this issue in some detail. A process referred to as "expert elicitation" is commonly used in risk analysis studies to determine a consensus risk tolerance level. As an analogous example, this kind of a process might be used to determine the acceptable level of risk that there is an exogenous termite in a shipment of imported lumber. In these kinds of cases, although it would be desirable to have the risk be as low as possible, setting the risk level too low typically introduces unacceptable consequences in other areas. To determine the difference between acceptable and unacceptable risk, it is helpful to set quantitative risk standards. Such risk standards have been used in martian PP for many decades for parameters like orbital lifetime and trajectory biasing. However, SR-SAG ultimately chose not to propose a consensus quantitative risk standard, for at least two reasons:

- The probability that T and a_w will exceed their threshold values in a given martian environment within 100 years of mission arrival cannot be quantitatively determined for all martian environments.
- Even if T and a_w exceed their threshold values in a specific martian environment, the probability of propagation of a population of mixed terrestrial microbes cannot be quantitatively estimated. There are too many additional factors involved that are poorly understood. We have minimal ability to do this even on Earth.

For these reasons, SR-SAG opted instead to use a qualitative determination of acceptability of risk (to this group) for different martian environments (Table 6). This risk acceptability includes the likelihood that the environment exists at all, or will exist within 100 years (*e.g.*, the probability that volcanic activity will commence at a specific site); the risk that, if the environment exists, propagation is possible (gullies exist, but can organisms propagate there within 100 years?); and the risk that an extreme hypothesis that is difficult to test and impossible to reject will turn out to be correct. One of the essential aspects of exploration is that there is an element of the unknown, which is why the exploration is justified in the first place. Although there is risk in encountering the

TABLE 6. LISTINGS OF MARS ENVIRONMENTS OF CONCERN AS POSSIBLE SPECIAL REGIONS

Classification of martian environments by their potential to exceed the threshold conditions in temperature and water activity for microbial propagation (within the boundary conditions of the analysis)

A. Observed features for which there is a significant (but still unknown) probability of association with modern liquid water	**B. Observed features for which there is a low, but non-zero, probability of a relationship to modern liquid water**	**C. Not known to exist, but if examples could be found, would have a high probability of association with modern liquid water**
• Recent gullies and gully-forming regions • "Pasted-on" mantle	• Low-latitude slope streaks • Low-latitude features hypothesized to be glaciers • Features hypothesized to be massive subsurface ice	• Volcanic environments young enough to retain heat • Impact environments young enough and large enough to retain heat • Modern outflow channels

unknown, we have learned from our experience on Earth that, as scientific knowledge progressively increases, it is possible to develop strategies for carrying out exploration activity in a relatively safe (though not risk-free) way.

It is important to point out that the composition of the SR-SAG team was chosen to be representative of the Mars scientific exploration community. The team is entirely composed of Mars scientists, who are intensely interested in the future exploration of Mars. Although the judgments embodied in Table 6 describe SR-SAG's comfort zone, there are clearly members of the community who are both more and less conservative than this consensus position. The same is true of other stakeholders in Mars exploration, and SR-SAG makes no representation regarding their views. Should a more conservative risk posture be desired, there are several implementation options available, including sterilizing all spacecraft sent to Mars, sterilizing spacecraft sent to certain places on Mars, placing certain places on Mars off-limits, and not flying spacecraft to Mars at all.

9B. Special Regions on Mars Within the Temporal and Spatial Limits of This Analysis

Environments that definitely exceed the threshold conditions. SR-SAG cannot identify any regions on Mars in which the threshold conditions for propagation are exceeded at the present time.

Environments of concern. Using the guidelines described in this document (DEFINITION #2), there are martian environments (based on understanding as of April 2006) that might exceed the threshold conditions for propagation of terrestrial organisms. They are summarized in Table 6. Of these, only the pasted-on mantle and the gully-forming regions are thought to be of high enough concern for exceeding the threshold conditions during the next 100 years that the SR-SAG believes they should be treated as special regions for the purposes of PP.

Gullies and pasted-on mantle of concern here were formed at the same latitudes and on slopes where insolation conditions might (though this, too, is highly uncertain and inconclusive) be conducive to their development. The relationship between gullies and pasted-on mantle is complex—there might be more than one type of each, and they may have asynchronous, or no, relationship. These are clearly an area for future research attention. The following relationships are worth noting:

- There is a very real possibility that the gullies were formed by the action of liquid water, and the possibility (though less-well documented and studied as of the present time) that "pasted-on" mantle provided that water (at least in some cases).
- The gullies and pasted-on mantle both occur at middle latitudes.

Map distribution. Based on the analysis above, several boundaries of primary significance to interpreting the possible presence of special regions on Mars can be shown in map format (Fig. 22) (for additional details on the construction of this map, see Appendix).

- **BOUNDARY A:** The 6 counts/s isopleth (using summer data only) from the GRS instrument (Fig. 5).
- **BOUNDARY B:** The most equatorward position of the limit of ice stability at a depth of 5 m, using the three planetary-scale thermodynamic models presented at the Mars Water Conference in February 2006. Each of these models uses somewhat different methodology and somewhat different input parameters, so the derived results are somewhat different (although they are all based on the same physics). The SR-SAG is not in a position to judge which, if any, of these models is correct. However, they all have the same general form, with an equatorial belt where none of them shows that ice is stable within 5 m of the surface, north and south polar zones where all of them agree there is continuous ice within 1 m of the surface, and an intermediate zone where ice is discontinuous or within 1–5 m of the martian surface. For the purpose of Fig. 22, the equatorial and polar zones were mapped only where all three models agreed; thus, the mid-latitude zone also incorporates all of the model-dependent uncertainty.
- **BOUNDARY C:** The observed distribution of gullies and mid-latitude mantles has a relatively well-defined equatorward limit (see Figs. 10 and 12). The map distribution of the roughened mantles of Mustard *et al*. (2001) is

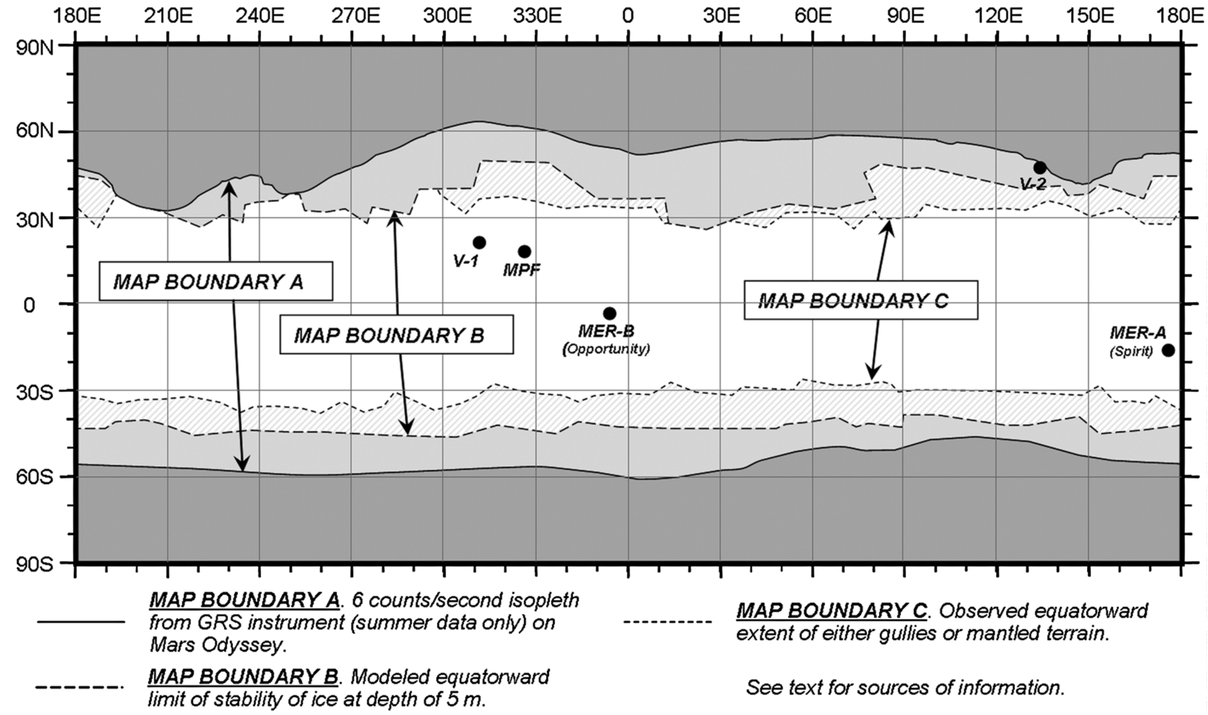

FIG. 22. Map of the stability of ice in the shallow martian subsurface, as shown by synthesis of results from three thermodynamic models presented at the 2006 Mars Water Conference by Aharonson and Shorghofer (2007), Mellon and Feldman (2006), and Chamberlain and Boynton (2006). (For base map details, see legend to Fig. 10.)

known (Fig. 12), but the "pasted-on" mantles of Christensen (2003) have not been sufficiently mapped. As a practical matter, the distribution of gullies and roughened mantles is very similar. Because of the possibility that either could be an environment where the threshold conditions for microbial propagation are exceeded, this boundary is drawn to encompass both.

The map areas between these boundaries lead to the following interpretations:

- The equatorial belt (*i.e.*, the region equatorward of all three map boundaries) is a region within which ice is thermodynamically unstable within 5 m of the martian surface at the present time. Gullies are absent. Although shallow liquid water could be present in a disequilibrium environment, such environments have yet been identified. The youngest volcanics on Mars are present within this region (see Fig. 17), but there is no evidence that they are young enough to retain enough heat to support modern liquid water. To date, despite ongoing searches, no thermal anomalies have been identified.
- Between Boundaries B and C, ice is thermodynamically unstable within 5 m of the martian surface at the present time, and gullies, pasted-on mantle, and mantles are locally present. The geologic processes associated with gullies and pasted-on mantle is still a subject of active research, but there is a significant possibility that they could be continuously or episodically active, and their activity could cause the threshold conditions for temperature and water activity to be exceeded. Because these features are only locally present within this region (especially on certain kinds of crater walls), there are large areas where there is no evidence of either gullies or mantles. Note that gullies and mantles are not limited in their map distribution to the area between Boundaries B and C—both extend to considerably higher latitude than Boundary B.
- Between Boundaries A and B is a region within which there is evidence (from models) that ice is stable at a depth from 1 to 5 m, or within which it is discontinuously or seasonally present. Such discontinuities can be caused by variations in albedo, variations in thermal inertia, slope effects (*e.g.*, poleward-facing slopes

get less sun that equatorward-facing slopes). The longitudinal variation in the position of Boundary B is primarily related to variations in thermal inertia of the martian near-surface materials. The worst-case crash scenarios within this region are potentially capable of penetrating to ice.

- The area poleward of Boundary A represents the polar regions, where the GRS data have been interpreted to show essentially continuous shallow frozen ground and ice caps at each pole. It is too cold for naturally occurring liquid water at any time during the year.

10. DISCUSSION OF SPACECRAFT-INDUCED SPECIAL REGIONS

The analysis and discussion to this point have focused on identification of naturally occurring environments on Mars within which terrestrial life forms might be able to propagate. To complete the analysis, consideration was given to whether the arrival of a spacecraft could induce a special region on Mars even when one did not exist at the landing site beforehand. Spacecraft are capable of generating heat and carrying liquid water, both of which could have an effect on the threshold conditions for propagation. Most importantly, in regions where there is (or may be) ice present, it is conceivable that local heating would result in the formation of liquid water for some amount of time. If conditions were created that exceeded the temperature and water activity thresholds for significant periods, replication of spacecraft-borne terrestrial bioburden could not be ruled out.

A series of representative scenarios for different mission types, including orbiters, landers/rovers, balloons, and drill missions were considered, for both nominal and non-nominal (crash) scenarios. Also considered were different operational factors: descent engine exhaust plumes, power/heat sources, roving into special regions, on-surface activities [*e.g.*, sampling (scoops, drills, rock abrasion tools, melt probes)] and burn-up/break-up scenarios for discarded descent hardware.

The complexity of considering all the potential scenarios with respect to induced special regions led to the conclusion that general categorization is not possible or practicable. The recommendation, therefore, is that analysis of spacecraft-induced special regions for Mars missions should be done on a case-by-case basis, with mission project teams being required to produce something equivalent to a "Mars environmental impact assessment" as part of the early stage mission planning. (Such an analysis may form part of the support documentation for the mission certification request.)

The use of radioisotope thermal generators on spacecraft will necessitate an extended analysis, since they can act as a perennial heat source, creating temperatures local to the radioisotope thermal generator above the threshold temperature condition. In the case of a non-nominal landing, resulting in the co-location of near-surface ice, an radioisotope thermal generator, and contaminated spacecraft parts, it would be possible to have a disequilibrium condition for many months, exceeding the proposed threshold conditions for propagation of terrestrial organisms.

A key factor in the assessment of spacecraft-induced environmental changes is the minimum duration for the conditions to be above the thresholds before qualifying as a special region and, therefore, a PP concern. At one extreme, the pressure or friction associated with sampling, for example with a scoop, may cause melting of ice in the time frame of seconds or less. At the other extreme, longer-duration increases in temperature associated with spacecraft activity can be predicted, for example, deliberate or accidental contact of a perennial heat source with the surface. In this scenario, the local ice will act as a heat sink, dissipating heat rapidly, but there may be a longer period of time (up to years) where the temperature and/or a_w thresholds are exceeded.

It is recognized that microbial propagation is not an instantaneous process—a finite amount of time above the temperature and a_w thresholds is necessary for the mechanics of biological propagation to take place. As an example, food spoilage in a refrigerator demonstrates that microbes are active at low temperatures but are slow growing, with the amount of time for spoilage dependent on temperature. Unfortunately, for the analysis of growth rates pertinent to spacecraft-induced special regions, there are not the same volume and quality of data to draw from as for the analysis of the temperature and a_w thresholds presented in Sec. 5 above, and while there are abundant data on growth responses of organisms to temperature, little is available in the range of interest to this problem. It was possible, however, to search the literature for the highest documented growth rates at the temperatures of interest and report here on the growth rates of whatever microorganism was the subject of that literature. Organisms isolated from polar environments seem to be the most cold tolerant that have been tested, whereas microorganisms responsible for food spoilage under refrigeration seem to *grow* the best at cold but not extreme temperatures. We found that the highest documented growth rates at the temperatures of interest (−15°C to +5°C) were based on an early study of *Sporotrichum carnis* (Haines, 1931; Ratkowsky *et al*., 1982). We concluded that significant replication of this or any other terrestrial microorganism would not occur if the temperature excursion to a maximum of −5°C did not exceed 22 h, to 0°C did not exceed 3 h, and to 5°C did not exceed 1 h. Subsequent work on growth rate models (Ratkowsky *et al*., 1983), based on estimates for the doubling time of terrestrial organisms in ideal culture at 5°C, support these conclusions. (For an analysis pertinent to conditions for possible microbial growth at Mars, it is worth noting that the boiling point of water at martian atmospheric pressure of 8.6 mbar is 5°C; therefore, there is no need to extend this kind of analysis to temperatures higher than shown in Table 7.) The proposed times in Table 7 may be overly restrictive in that they are based on organisms in the exponential growth phase and do not take into account the latency that would precede growth and division once adequate temperatures were attained.

It was further concluded that a cumulative limit for qualification as a special region could *not* reasonably be set. While some organisms are able to "bank" metabolic activity and benefit from serial exposure to favorable conditions, few or no data are available with regard to evaluating the affect of freeze thaw cycles that range from +5°C to −20°C or below on microbial metabolism. Indeed, it is reasonable to expect that repeated freeze/thaw cycles would have cumulative negative effects on the majority of organisms; the SR-SAG was not able to translate this into a quantitative implementation guideline that is based on experimental data. Further conservatism is present in this recommendation since over this period of time the other stressors in the martian environment would mitigate against faster replication. Microbial psychrophily is an active area of research, and as additional studies targeting physiological responses and cold adaptation provide new information, future assessments can define additional boundaries, including cumulative limits.

FINDING. It is possible for spacecraft to induce conditions that could exceed for some time the threshold conditions for biological propagation, even when the ambient conditions were in equilibrium before the spacecraft arrived. Whether a special region is induced or not depends on the configuration of the spacecraft, where it is sent, and what it does. This possibility is best evaluated on a case-by-case basis.

11. APPENDIX (DERIVATION OF FIG. 22)

Using the following process, a map of the martian shallow equilibrium ice was developed:

- The data from the GRS instrument on Mars Odyssey are widely accepted as a clear indication of high-latitude shallow ground ice. For the purpose of mapping, only summer data from both hemispheres are used (winter CO_2 frost obscures the ice signature by adding hydrogen-poor mass atop the soil—seasonal CO_2 can be as much as a meter or more at high latitudes). Using an arbitrarily selected threshold

TABLE 7. PROPOSED TIMES FOR WHICH LOCALIZED SPACECRAFT-INDUCED ENVIRONMENTS MAY EXCEED THE TEMPERATURE THRESHOLD OF −20°C WITH NO CELL REPLICATION RESULTING

Not-to-exceed temperature of spacecraft-induced environment	*Elapsed time before replication of terrestrial organism could occur*
−5°C	22 h
0°C	3 h
5°C	1 h

value of 6 counts/s, the region of permanent shallow ground ice can be shown. Note that poleward of that boundary, there are fewer than 6 counts/s (the epithermal neutrons go down as hydrogen goes up) and that the position of the 6 counts/s threshold value should be considered blurred on a 600 km scale because of the GRS neutron footprint.

- Several equilibrium thermodynamic models for Mars have been calculated at a planetary scale. These models predict the distribution of ice equatorward of the GRS 6 counts/s threshold. Because the GRS cannot detect deeper than about a meter, this gives us a way to model the ice distribution to deeper depths. At the 2006 Mars Water Conference (http://es.ucsc.edu/~fnimmo/website/mars2006.html), the models developed by three independent research teams (Chamberlain and Boynton, 2006; Mellon and Feldman, 2006; Aharonson and Schorghofer, 2007) were presented. These models use somewhat different methodology and inputs, but the results have a very similar structure. In each case, there is a north and south mid-latitude belt that can be thought of as discontinuous ice, and/or a zone of ice less detectable by GRS because of its depth, along with an equatorial belt with no near-surface ice. The position of the mid-latitude belts is somewhat different in the three models. To allow for appropriate conservatism, a boundary was drawn (Boundary B on Fig. 22) that encompasses the most equatorward indication of ice in any of the models. Thus, Boundary B also incorporates model-dependent uncertainty.
- Equilibrium thermodynamic models show that the depth to the top of the ice table increases abruptly at about the position of the dashed line. This has been studied extensively (*e.g.*, Farmer and Doms, 1979; Paige, 1992). It is typical in model results for the transition from 5 m to infinite to occur in less than 1° of latitude. In order to further represent appropriate conservatism, the north and south dashed lines were each shifted 1° of latitude towards the equator, so as to encompass possible ice within 5 m of the surface.

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13. ABBREVIATIONS

a_w, water activity; COSPAR, Committee on Space Research; GRS, Gamma Ray Spectrometer; HRSC, High Resolution Stereo Camera; Ma, age in millions of years; MEPAG, Mars Exploration Program Analysis Group; MER, Mars Exploration Rover; MGS, Mars Global Surveyor; MOC, Mars Orbiter Camera; MOLA, Mars Orbital Laser Altimeter; NASA, National Aeronautics and Space Administration; NRC, National Research Council; PP, planetary protection; SR-SAG, Special Regions–Science Advisory Group; THEMIS, Thermal Emissions Imaging System; UV, ultraviolet.

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